

# On the onset of galactic winds in quiescent star forming galaxies

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## ABSTRACT

**Context.** The hierarchical model of galaxy formation, despite its many successes, still overpredicts the baryons fraction locked in galaxies as a condensed phase. The efficiency of supernovae feedback, proposed a long time ago as a possible solution for this so-called “overcooling” problem, is still under debate, mainly because modelling supernovae explosions within a turbulent interstellar medium, while capturing realistic large scale flows around the galaxy is a very demanding task.

**Aims.** Our goal is to study the effect of supernovae feedback on a disk galaxy, taking into account the impact of infalling gas on both the star formation history and the corresponding outflow structure, the apparition of a supernovae-driven wind being highly sensitive to the halo mass, the galaxy spin and the star formation efficiency.

**Methods.** We model our galaxies as cooling and collapsing NFW spheres. The dark matter component is modelled as a static external potential, while the baryon component is described by the Euler equations using the AMR code RAMSES. Metal-dependent cooling and supernovae-heating are also implemented using state-of-the-art recipes coming from cosmological simulations. We allow for 3 parameters to vary: the halo circular velocity, the spin parameter and the star formation efficiency.

**Results.** We found that the ram pressure of infalling material is the key factor limiting the apparition of galactic winds. We obtain a very low feedback efficiency, with supernovae to wind energy conversion factor around one percent, so that only low circular velocity galaxies give rise to strong winds. For massive galaxies, we obtain a galactic fountain, for which we discuss the observational properties.

**Conclusions.** We conclude that for quiescent isolated galaxies, galactic winds appear only in very low mass systems. Although that can quite efficiently enrich the IGM with metals, they don't carry away enough cold material to solve the overcooling problem.

**Key words.** galaxies: formation – galaxies: evolution – galaxies: structure – methods: numerical

## 1. Introduction

The hierarchical model of structure formation, for which massive galaxies grow by mergers of smaller satellites or by filamentary accretion of gas, is now a well-established theory (the so-called CDM paradigm), that compares favorably with observations. In this standard picture, dark matter is the main component that drives structure formation via gravitational instability from primordial density perturbations. On the other hand, numerical simulations including gas dynamics with radiative cooling and some recipes of star formation tend to produce galactic disks that are too small, too dense, with too low angular momentum and too many condensed baryons, when compared to observations. These various problems, called “overcooling” (Dekel & Silk 1986; Blanchard et al. 1992; Navarro & White 1993; Yepes et al. 1997; Gnedin 1998; Hultman & Pharasyn 1999; Somerville & Primack 1999; Cole et al. 2000; Kay et al. 2002; Springel & Hernquist 2003) or “angular momentum” (Navarro & Benz 1991; Sommer-Larsen et al. 1999; Steinmetz & Navarro 1999;

Bullock et al. 2001; Maller & Dekel 2002; Abadi et al. 2003; Sommer-Larsen et al. 2003; Read & Gilmore 2005) problem, may all have the same physical origin. Nevertheless, they are considered as weak points in the CDM theory.

Supernovae-driven winds are a key ingredient of current galaxy formation models, in order to suppress the formation of low-mass galaxies and maybe to solve the “overcooling” problem. They can potentially suppress the fast consumption of gas by transferring cold and dense gas into the hot and diffuse surrounding medium of a galaxy. Thus, a fraction of gas would not collapse into cold dense molecular clouds and would not transform into stars. That is the most favored explanation for the disruption of star formation and the lack of baryons available in stars (Springel & Hernquist 2003; Rasera & Teyssier 2006; Stinson et al. 2006). Feedback is also invoked to reproduce the large disc galaxies observed at our epoch (Weil et al. 1998; Thacker & Couchman 2001), the morphology of galactic discs (Sommer-Larsen et al. 2003; Okamoto et al. 2005), to stabilize the gaseous disc against the Toomre instability (Robertson et al. 2004) or to fit the temperature of the intergalactic medium (IGM) at redshift  $\sim 3$  (Cen & Bryan 2001).

The proper modelling of galactic winds is quite difficult to handle, especially in cosmological simulations, because of the very large scale separation between the galactic outflow in the IGM and the supernova explosions within the interstellar medium (ISM). But it represents an unavoidable task in order to solve the weakest points of the hierarchical model, both in semi-analytical models (White & Frenk 1991; Somerville & Primack 1999; Kauffmann et al. 1999; Cole et al. 2000; Hatton et al. 2003; Monaco 2004; Bertone et al. 2005) and in numerical simulation (Cen & Ostriker 1992; Navarro & White 1993; Mihos & Hernquist 1994; Katz et al. 1996; Thacker & Couchman 2000; Scannapieco et al. 2001; Springel & Hernquist 2003; Rasera & Teyssier 2006) of galaxy formation. Observational evidence for galactic outflows have already been pointed out by several authors (Bland & Tully 1988; Heckman et al. 1990; Ellison et al. 2000; Heckman et al. 2000; Pettini et al. 2001, 2002; Adelberger et al. 2003; Bouché et al. 2006). They are usually associated to massive starbursts, for which very strong outflows are reported: for one solar mass of stars formed in the galaxy, between 1 and 5 solar mass of gas are ejected in the winds of starburst dwarf galaxies (Martin 1999).

Despite the difficulties of modelling supernovae explosions within a turbulent, multiphase and magnetized ISM, understanding the physics of the resulting large scale outflows is also a challenge. As explored by Fujita et al. (2004) in the context of an isolated, pre-formed galactic disc, the ram-pressure of infalling material might be the main limiting factor for galactic winds to exist. The gravitational potential is likely to play a minor role in the development of galactic winds, especially in low-mass galaxies where the escape velocity is far lower than the velocity of the wind. This fact, outlined first by Fujita et al. (2004), is in contradiction with most semi-analytical studies so far, for which only the escape velocity of the parent halo determines if a wind can develop or not. In this paper, using a self-consistent modelling of gas cooling and accretion from the hot surrounding medium into the disc, we will demonstrate that it is indeed the ram pressure of the infalling gas that matters for galactic outflows.

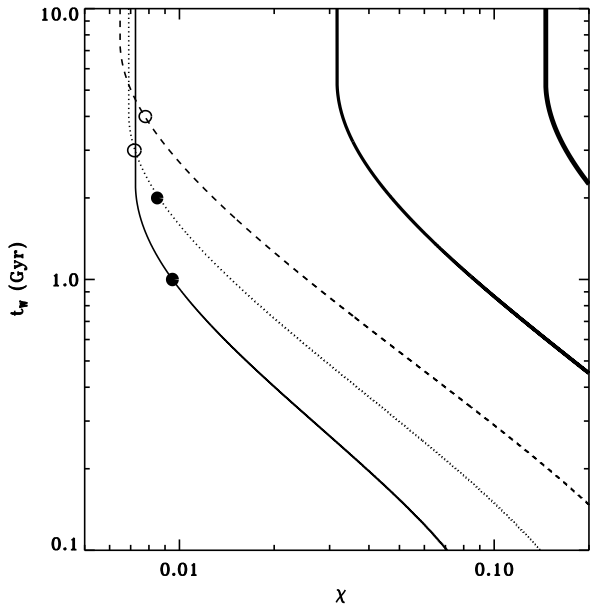
To illustrate further the difficulties in modelling supernovae feedback, we note that the mass ejection rates measured in galactic winds obtained in numerical simulations are widely discrepant. Some authors report very efficient winds, even in the quite massive galaxies (Springel & Hernquist 2003; Tasker & Bryan 2006), while other authors obtained very low ejection efficiency, mainly in dwarf galaxies (Mac Low & Ferrara 1999). We believe that these differences arise partly from the different numerical technics used, and partly from the different boundary conditions imposed around the galaxy. In this paper, we report the first self-consistent adaptive mesh refinement (AMR) simulation of a collapsing NFW (Navarro et al. 1996) spherical cloud forming a galactic disk from the inside out. In the recent literature, this same problem was addressed using smooth particle hydrodynamics (SPH) technics by Springel & Hernquist (2003). Grid-based simulations were also performed, but using isolated, pre-formed rotating discs (Fujita et al. 2004; Tasker & Bryan 2006). One of

the key outcome of such numerical simulation is the fraction of the injected supernovae energy which is transferred up to the large scale galactic wind and is truly available for feedback (Fujita et al. 2004; Scannapieco et al. 2006).

We are indeed facing a major difficulty when introducing feedback in numerical simulations that include radiative losses: the characteristic time scale of radiative cooling in dense regions where star formation occurs is very fast, usually faster than the simulation time step, so that one relies on implicit time integration schemes. As a consequence, if one handles supernovae heating by dumping thermal energy in the same dense regions, most (if not all) of the energy is radiated away before any large scale flow develops. This spurious effect has been identified for a long time, both in cosmological simulations (Katz 1992, Mihos & Hernquist 1994, Katz et al. 1996) and in interstellar medium simulations (Wada et al. 2000, Joung & Mac Low 2006). We follow the recipe proposed by Mori et al. (1997) and Gnedin (1998) by adding around each supernova event density, velocity and energy profiles based on the Sedov blast wave solution, in order to get the correct kinetic to thermal energy fraction in the flow. In this way, a large fraction of kinetic energy is injected into the disc, without being radiated away on the spot. This will translate into a turbulent gaseous disc, and in some cases into a large scale outflow. To account for the small scale effect of supernovae feedback, we follow the multiphase approach of Yepes et al. 1997, Ascasibar et al. 2002, Springel & Hernquist 2003, and Marri & White 2003, for which the supernovae thermal energy is collectively accounted for as a new equation of state for the dense ISM, depending on the gas local density (polytropic equation of state). In this way, we can reproduce a stiffer equation of state for the dense gas, leading to thicker, more stable discs (Toomre 1964, Robertson et al. 2004). Another problem arises because the scales where radiative losses of SNe remnants occur are largely unresolved ( $\lesssim 1$  pc). It is therefore likely that the energy transfer of SN remnants to the ISM is much lower (Larson 1974; Thornton et al. 1998) than the energy transferred from superbubbles initiated by multiple localized SNe (Dekel & Silk 1986; Melioli & de Gouveia Dal Pino 2004). In reality, modeling properly the energy deposition requires to resolve the coherence length of the SNe explosions. Thus, we have to bear in mind that our superbubble model entails many assumptions.

Within this framework, we would like to answer the following questions: What is the conversion efficiency between the small scale supernovae luminosity and the large scale outflow luminosity? What are the conditions for a galactic wind to develop and escape from the parent halo potential well? What is the mass ejection rate of such a wind? What is the metallicity of the wind and other associated observational signatures?

Although the cooling NFW halo approach that we present in this paper represents an important improvement over previous isolated disc grid-based simulations, we are still a long way from realistic cosmological simulations. Such idealized simulations are nevertheless useful to understand the physical processes responsible for the apparition of a galactic wind subject to boundary conditions which are relevant for cosmology. They also give us the possibility to derive a simple analytical model



**Fig. 1.** The epoch of wind formation  $t_W$  in Gyr as a function of the hydrodynamical efficiency  $\chi$ . We have assumed three different average star formation time  $\langle t_* \rangle = 1$  Gyr (thin solid line),  $\langle t_* \rangle = 2$  Gyr (dotted line) and  $\langle t_* \rangle = 4$  Gyr (dashed line) for the  $10^{10} M_\odot$  halo, and the same  $\langle t_* \rangle = 3$  Gyr for the  $10^{11} M_\odot$  (solid line) and the  $10^{12} M_\odot$  (thick solid line) halos. Open circles are for Sc ( $t_W = 3$  Gyr) and Sd ( $t_W = 4$  Gyr), circles are for Sa ( $t_W = 1$  Gyr) and Sb ( $t_W = 2$  Gyr).

that highlights the importance of infall in the process of wind formation. Our paper is organized as follows. In section 2 we present our analytical model and introduce the relevant quantities for constraining the galactic wind formation epoch. In section 3, we present our simulation settings, with the physics involved in our different disc simulations, as well as some numerical details. In section 4, we present our main result, namely the formation of either a galactic outflow or a galactic fountain. We compare the actual wind formation epoch with our analytical prediction. We also describe in details the flow structure in each case, and discuss its observational signature.

## 2. Galactic winds from cooling NFW halos

In this section, we will describe the simplified problem we would like to study using high-resolution numerical simulations, as well as a simple analytical model that will guide us in interpreting the results. We model an isolated NFW halo (Navarro et al. (1996)), with initially similar distribution for the gas and the dark matter components (Ascasibar et al. 2003). We follow the initial settings proposed by Springel & Hernquist (2003), so our work can be directly compared to them. Imposing identical profiles for the gas and for the dark matter is especially important in the outer part, in order to have realistic accretion rates at late times, when the

outer region of the halo are falling into the inner disc. Both density profiles are therefore given by

$$\rho = \frac{\rho_s}{r_s(1+r/r_s)^2}, \quad (1)$$

so that the total integrated mass is

$$M(< r) = 4\pi\rho_s r_s^3 \left( \ln(1+r/r_s) - \frac{r/r_s}{1+r/r_s} \right), \quad (2)$$

$\rho_s$  is the characteristic density of the halo and  $r_s$  its characteristic radius. We define here the Virial radius as the radius where the average density is equal to 200 times the *critical* density, assuming  $H_0 = 70$  km/s/Mpc. We consider a constant gas fraction of 15% throughout the halo. The temperature profile is adjusted so that a strict hydrostatic equilibrium of the gas sphere is maintained in the non-cooling, non-rotating case. The halo is truncated at 2 Virial radii: outside the halo, we imposed a constant gas density equal to a small fraction ( $10^{-4}$ ) of the density at this radius. Note that this truncation radius is a key parameter in this study. The size of the box is taken equal to 6 Virial radius. The concentration parameter  $c = r_{vir}/r_s$  of the NFW profile is taken to be 10, independant of the halo mass. Initially, our halo is slowly rotating, with an angular momentum profile corresponding to the average profile found in 3D cosmological simulations (Bullock et al. 2001), for which the specific angular momentum is

$$j(r) = j_{max} \frac{M(< r)}{M_{vir}} \quad (3)$$

with a spin parameter  $\lambda = J|E|^{1/2}/(GM_{vir}^{5/2})$ , for which we consider only two representative values  $\lambda = 0.04$  and  $0.1$ . Starting from this initial equilibrium configuration, the gas is allowed to radiate its thermal energy using standard, *metal-dependant* cooling processes. Note that initially we assume a constant metallicity equal to  $10^{-3}$  solar throughout the halo. In agreement with the standard picture of galaxy formation, a centrifugally supported disc quickly forms in the halo centre, gradually growing from the inside out. Its size depends strongly on the chosen spin parameter. When enough dense material is formed within the disc, star formation is allowed to proceed in a quiescent way, using a star formation recipe based on a standard Schmidt law (see Eq. 10). According to Salpeter (1955) IMF, for each solar mass of stars formed,  $\eta_{SN} \approx 0.1$  solar masses are recycled into supernovae ejecta, and drive strong shock waves into the surrounding ISM material.

Our goal is to implement these two rather standard physical ingredients in the RAMSES code, in order to study the formation of galactic winds. This will be described in the next section. In the present section, we would like to develop a simple analytical model that allows us to capture the basic mechanism driving galactic winds. Our picture contrasts with the one discussed in Fujita et al. (2004), who assumed a pre-formed gaseous galactic disc, with no self-consistent star formation. A blast wave was initiated at the very center of the disc, with an energy budget consistent with massive star bursts. Depending on the blast wave and disc parameters, two possible scenarii emerge: “blow out” for which the disc is completely destroyed, and “blow away”, corresponding to a galactic wind escaping

the galactic disc. These authors have demonstrated that the main limiting factor for a wind to appear in the “blow away” case is the amount of infalling gas, coming from the surrounding halo. In the present study, we will address this question using a different approach, based on quiescent star formation and infall from a cooling NFW halo.

Let us consider that the gas is cooling instantaneously, so that pressure support is lost from the very beginning. Each spherical shell is therefore free-falling, gradually accelerating towards the center. We compute the accretion rate of halo material onto the disc using a standard semi-analytical recipe (e.g. Hatton et al. 2003), namely by taking the time derivative of the integrated accreted mass

$$M_{acc}(< t) = M(< r_{ff}(t)), \quad (4)$$

where  $r_{ff}(t)$  is the radius for which the free-fall time  $t_{ff}(r)$  is equal to the current time  $t$ . The gaseous disc mass therefore grows by accretion and decays because of star formation, leading to a simple differential equation for its time evolution

$$\dot{M}_g = \dot{M}_{acc} - \frac{M_g}{\langle t_* \rangle}. \quad (5)$$

Here,  $\langle t_* \rangle$  is the (mass-weighted) average star formation rate within the disc. It depends on the exact disc geometry (size and thickness) and on the Schmidt law parameters. As a first order approximation, one can show that

$$\langle t_* \rangle \propto t_0 \lambda_0 : \quad (6)$$

small discs form stars more efficiently than large ones. If one assumes that  $\langle t_* \rangle$  remains constant in time, one gets the following formal solution

$$M_g(t) = \int_0^t \dot{M}_{acc}(u) \exp\left(\frac{u-t}{\langle t_* \rangle}\right) du. \quad (7)$$

We finally deduce the total supernova luminosity in the disc using

$$L_{SN} = \eta_{SN} \epsilon_{SN} \dot{M}_* = \eta_{SN} \epsilon_{SN} \frac{M_g}{\langle t \rangle_*}, \quad (8)$$

where  $\epsilon_{SN} \simeq 50$  keV is the specific energy produced by one single supernova ( $10^{51}$  erg for an average progenitor mass of  $10 M_\odot$ ). Part of the supernova energy contributes to sustain a turbulent multiphase ISM. Only a (small) fraction of this energy is converted into actual wind luminosity escaping from the dense gaseous disc. The fraction of energy that manage to escape from the dense gaseous disc will depend on the disc characteristic (thickness, size, gas content). In our simulations, these properties will be specified by the the spin parameter of the halo and by the star formation time scale of the Schmidt law. We parametrize this unknown conversion factor  $\chi$  by defining the wind luminosity as  $L_W = \chi L_{SN}$ . Our goal is to perform numerical simulations of star forming discs in order to compute the actual conversion efficiency resulting from various, complex hydrodynamical processes within the disc. We thus call this parameter  $\chi$  the *hydrodynamical efficiency*.

Our goal is now to derive under which conditions a galactic wind can form within the halo. We know from previous studies

(Fujita et al. 2004) that, in presence of infalling gas, the main factor preventing the wind from blowing out is the ram pressure of the surrounding material. This is a much more stringent constraint than any criterion based on the escape velocity. If we assume that the free-falling cold gas hits the disc plane with terminal velocity, we can compute the accretion luminosity as

$$L_{acc}(t) = \frac{1}{2} V_0^2(r_{ff}(t)) \dot{M}_{acc} \quad \text{and} \quad V_0^2(r) = -2 \int_0^r \frac{GM(x)}{x^2} dx \quad (9)$$

We define the *wind break-out epoch*  $t_W$  as the time for which the wind luminosity exceeds the accretion luminosity  $L_{acc}(t_W) = L_W(t_W)$ . In Figure 1, the wind break-out epoch is shown for various halo masses, as a function of our unknown parameter  $\chi$ . One clearly sees that the smaller the hydrodynamical efficiency, the later the wind will blow out of the disc. More importantly, *for each halo mass, there is a minimum efficiency below which no wind can form at all*. This minimal value for  $\chi$  is around 0.7% (resp. 3% and 15%) for a  $10^{10} M_\odot$  halo (resp.  $10^{11} M_\odot$  and  $10^{12} M_\odot$ ). We conclude from this very simple toy model that it is critical to determine the actual hydrodynamical efficiency in quiescent star forming galaxies. Our goal is here to perform high resolution numerical simulations of isolated galaxies, to measure the epoch when galactic winds appear and deduce from  $t_W$  the corresponding (most likely) hydrodynamical efficiency.

### 3. Numerical methods

Our simulations were performed with the Adaptive Mesh Refinement (AMR) code RAMSES (Teyssier 2002). The gas evolution is computed using a second-order Godunov scheme for the Euler equations, while collisionless star particles' trajectories are computed using a Particle-Mesh solver. The dark matter component is accounted for as a constant background gravitational potential. Gas cooling is taken into account as a source term in the energy equation. The cooling function is computing using the Sutherland & Dopita (1993) cooling model, using a look-up table in the temperature and metallicity plane. Metals created in the surrounding of supernova explosions are advected by the hydrodynamics solver as a passive scalar.

We now describe in more details our 3 main physical ingredients used for the present study, namely star formation, supernova thermal feedback and supernova kinetic feedback.

#### 3.1. Star formation recipe

In each cell, gas is converted into star particles following a Schmidt law

$$\dot{\rho}_* = -\frac{\rho}{t_*} \text{ if } \rho > \rho_0, \quad \dot{\rho}_* = 0 \text{ otherwise}, \quad (10)$$

where  $\rho_0$  is an arbitrary density threshold defining what we consider here as being “interstellar medium”. The star formation time scale is proportional to the local free-fall time,

$$t_* = t_0 \left( \frac{\rho}{\rho_0} \right)^{-1/2}. \quad (11)$$



The two parameters  $\rho_0$  and  $t_0$  are poorly known and scale dependent. One standard approach consists in calibrating these numbers to the observed star formation rate in local galaxies, the so-called Kennicutt (1998) law, which translates roughly into  $\rho_0 \simeq 0.1 \text{ H.cm}^{-3}$  and  $t_0 \simeq 1 - 10 \text{ Gyr}$ . In the present study, we allow  $t_0$  to take 2 values, namely 3 or 8 Gyr corresponding to star formation efficiencies of respectively  $\sim 0.05$  and 0.02, last value is compatible with Krumholz & Tan (2007) if we suppose that star formation efficiency remains the same at very high ( $> 10^5 \text{ H.cm}^{-3}$ ) and unresolved densities (maximum density in our simulations stalls to a few  $10^1 \text{ H.cm}^{-3}$  and  $10^2 \text{ H.cm}^{-3}$  respectively for the  $10^{10} M_\odot$  and  $10^{11} M_\odot$  due to polytropic equation of state). When a gas cell is eligible for star formation,  $N$  collisionless star particles are spawned using a Poisson random process, with probability

$$P(N) = \frac{\lambda_P}{N!} \exp(-\lambda_P), \quad (12)$$

where the mean value is

$$\lambda_P = \left( \frac{\rho \Delta x^3}{m_*} \right) \frac{\Delta t}{t_*}. \quad (13)$$

The mass of a star particle is taken to be an integer multiple of  $m_* = \rho_0 \Delta x_{min}^3 / (1 + \eta_{SN} + \eta_W)$  (see Appendix A for details on  $\eta_W$ ), and is therefore directly connected to the code spatial resolution (Rasera & Teyssier 2006). The minimum mass of a single star particle is therefore  $\sim 500 M_\odot$ , but high density regions can spawn more massive star particles. We also ensure that no more than 90% of the gas in a cell is depleted by the star formation process.

### 3.2. Polytropic equation of state

The second important ingredient for our galactic disc simulations is the thermal feedback of supernovae into the ISM, based on the multiphase model of Ascasibar et al. (2002) and Springel & Hernquist (2003). The idea is to assume that the ISM is driven by small scale effects (turbulence, thermal instability, thermal conduction, molecular cloud formation and evaporation) that quickly reach steady-state and lead to a quasi-equilibrium thermal state, for which the average temperature is a function of the mass density alone. In practice, when the gas density exceeds the ISM threshold  $\rho_0$ , the gas temperature is forced to be equal to or higher than

$$T = T_0 \left( \frac{\rho}{\rho_0} \right)^{\gamma_0 - 1} \quad (14)$$

where  $\gamma_0$  is the polytropic index, and can be chosen to the exact value of the multiphase model (Springel & Hernquist 2003) or to a constant value for sake of simplicity. If  $\gamma_0$  is chosen between 4/3 et 2, as prescribed by Springel & Hernquist (2003), the resulting equation of state gives similar results for the disc thickness and maximum density. Our standard choice is here  $\gamma_0 = 5/3$ .

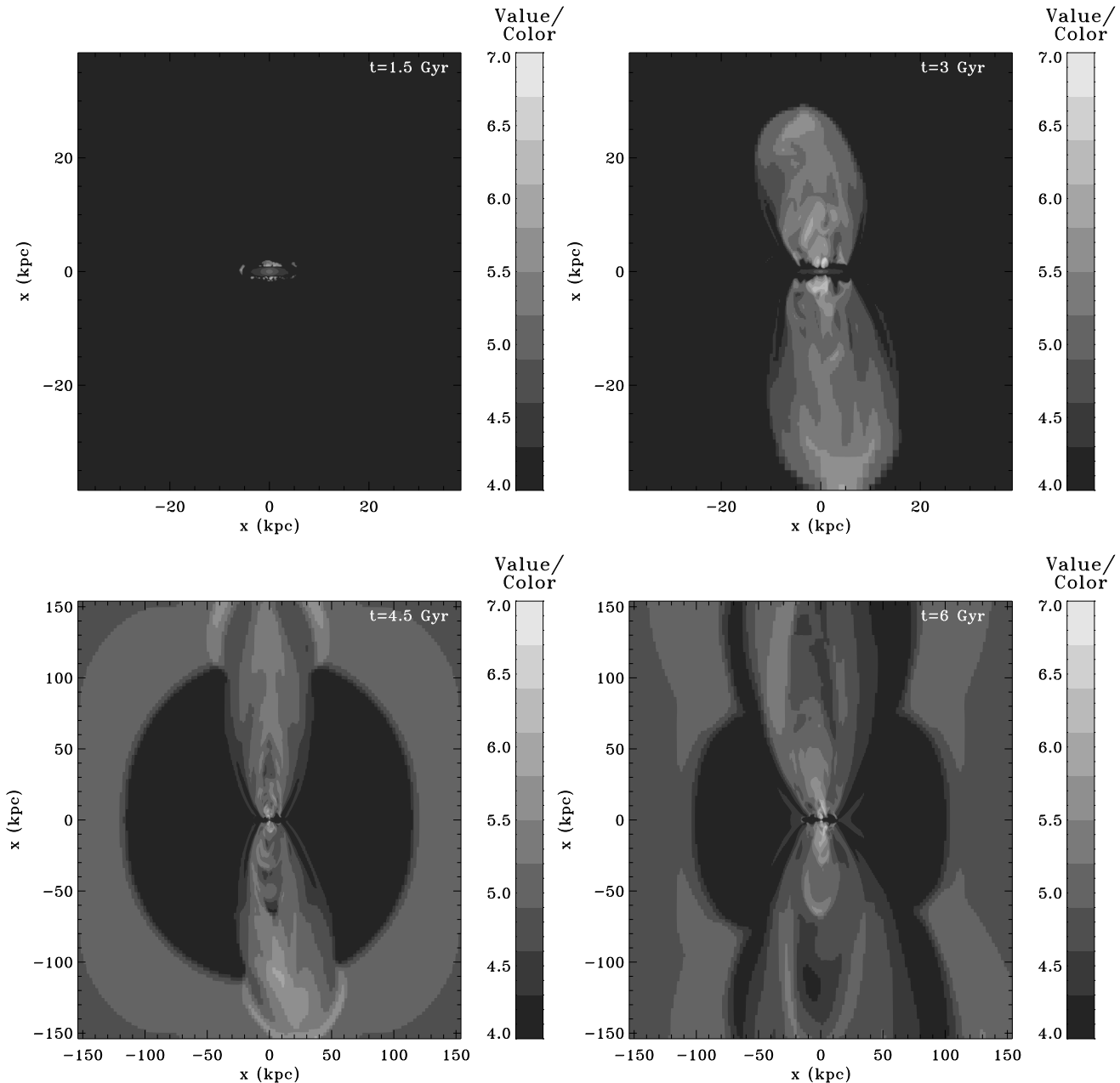
### 3.3. Blast wave model

The main problem of the previous thermodynamical approach is that feedback is only accounted for at mesoscopic scales. On

the other hand, it is now an observational fact that very large supernovae-driven bubbles appear quite naturally in star forming galaxies. These superbubbles are built by the coalescence of numerous supernovae remnants, who collectively create these large-scale features, driving very large, macroscopic flows that eventually give rise to galactic winds or galactic fountains. Current cosmological simulations do not have the necessary resolution to resolve individual supernovae blast wave. On the other hand, they do have the resolution to resolve superbubbles. Our goal is here to inject kinetic energy in the form of spherical blast waves of size comparable to galactic superbubbles, namely  $r_{SN} = 100$  to 200 pc. This kinetic energy injection can be considered as a turbulent forcing term with injection scale  $r_{SN}$ .

It is now well-established that supernovae energy can not be released only by direct thermal energy injection (Navarro & White 1993). For high density star forming region, the gas radiates away all this thermal energy in one time step due to very fast atomic cooling. The consequence is that SNe have no effect on the dynamics. Various methods have been proposed to correctly incorporate the SNII kinetic feedback into numerical simulations. The first idea, discussed for example in Governato et al. (2006) is to artificially stop radiative cooling in the region where the SNII explosion occurs, for a time long enough for the blast wave to develop and expand. This trick allows for a large enough fraction of the SNII energy to be converted into actual gas kinetic energy. The second approach is to directly inject kinetic energy into the surrounding gas. Springel & Hernquist (2003) used for that purpose high-velocity collisionless particles that travel over a rather large distance (up to a few kpc) before being re-incorporated in the fluid as new SPH particles.

In this paper, we directly add to each flow variable (density, momentum and total energy) a spherical blast wave solution of radius  $r_{SN} = 150$  pc, assuming that each supernovae releases half of its energy in this kinetic form (the other half is accounted for the polytropic equation of state). As soon as this distance is large enough compared to the grid resolution, we minimize spurious energy losses. Each time a “star” particle is created, we simultaneously remove an equivalent gas mass from the cell that is entrained by the SNII blast wave. Recalling that  $m_*$  is the mass locked into long-lived stars, the total mass removed from each star-forming cell is  $m_*(1 + \eta_{SN} + \eta_W)$ , where the entrainment parameter  $\eta_W$  is defined as in Springel & Hernquist (2003). We have used  $\eta_W \simeq 1$  in order to reproduce roughly a Sedov blast wave solution in typical galactic discs for which the density is slightly greater to  $\rho_0$ . Note that in order to implement this numerical scheme, we have used “debris” particles to describe the Sedov blast wave profile around each SNII explosion. Details are given in the Appendix. The most important parameter is  $r_{SN}$ , the radius of the spherical blast wave. This sets the injection scale of the turbulent cascade in the disc (Joung & Mac Low 2006). This value, comparable to the typical size of galactic superbubbles (McKee & Ostriker 1977), is however larger than individual supernovae remnants, whose size just before the snow plow phase varies from 10 to 50 pc, depending on the exact early shock wave dynamics, whether adiabatic or evaporative



**Fig. 2.** Cut of the gas temperature in the Oyz plane for the Sd simulation at different epochs. The two top panels are a zoom 4 times of the simulation box and the two bottom panels are the entire simulation box. The colour scale gives the temperature in log(K). Note that length scales are not the same in each panel.

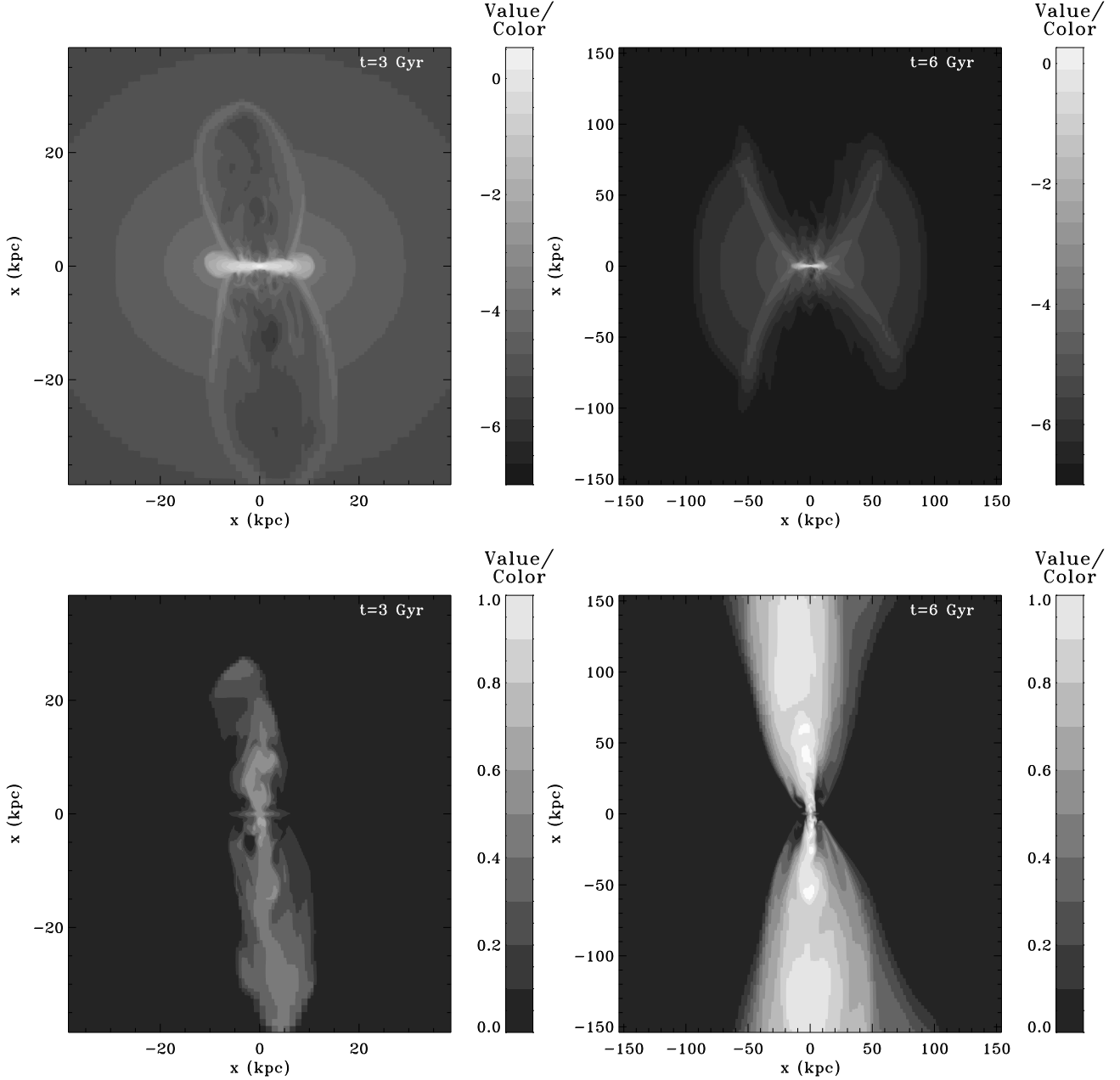
(Cioffi et al. 1988). We have not tried to study the impact of this important parameter on our results, in order to avoid clouding the purpose of this paper. We rather keep this parameter fixed to its fiducial value of  $r_{SN} = 150$  pc (McKee & Ostriker 1977, Shull & Silk 1979).

### 3.4. Simulation parameters

We use for the fluid solver simple “outflow” boundary conditions, for which all flow variable are assumed to have zero gradient at the boundary. All simulations are initialized with a NFW gas and dark matter halo in hydrostatic equilibrium, with

the same density profile for both fluids. The dark matter component is not explicitly simulated here; only baryons (stars and gas) are self-gravitating. We therefore add a fixed analytical gravitational potential to the solution of the Poisson equation.

We have performed our simulations for 2 different halo circular velocity, namely  $V_{vir} = 35$  and  $75$  km/s, which correspond to Virial masses of  $10^{10} M_{\odot}$  and  $10^{11} M_{\odot}$  at redshift zero. We have explored two different star formation efficiency, with timescales  $t_0 = 3$  Gyr and  $t_0 = 8$  Gyr. We use a coarse grid of  $128^3$  cells, in order to properly describe the large scale dynamics within the halo, which corresponds, in the RAMSES terminology, to a minimum level of refinement  $\ell_{min} = 7$ . The grid

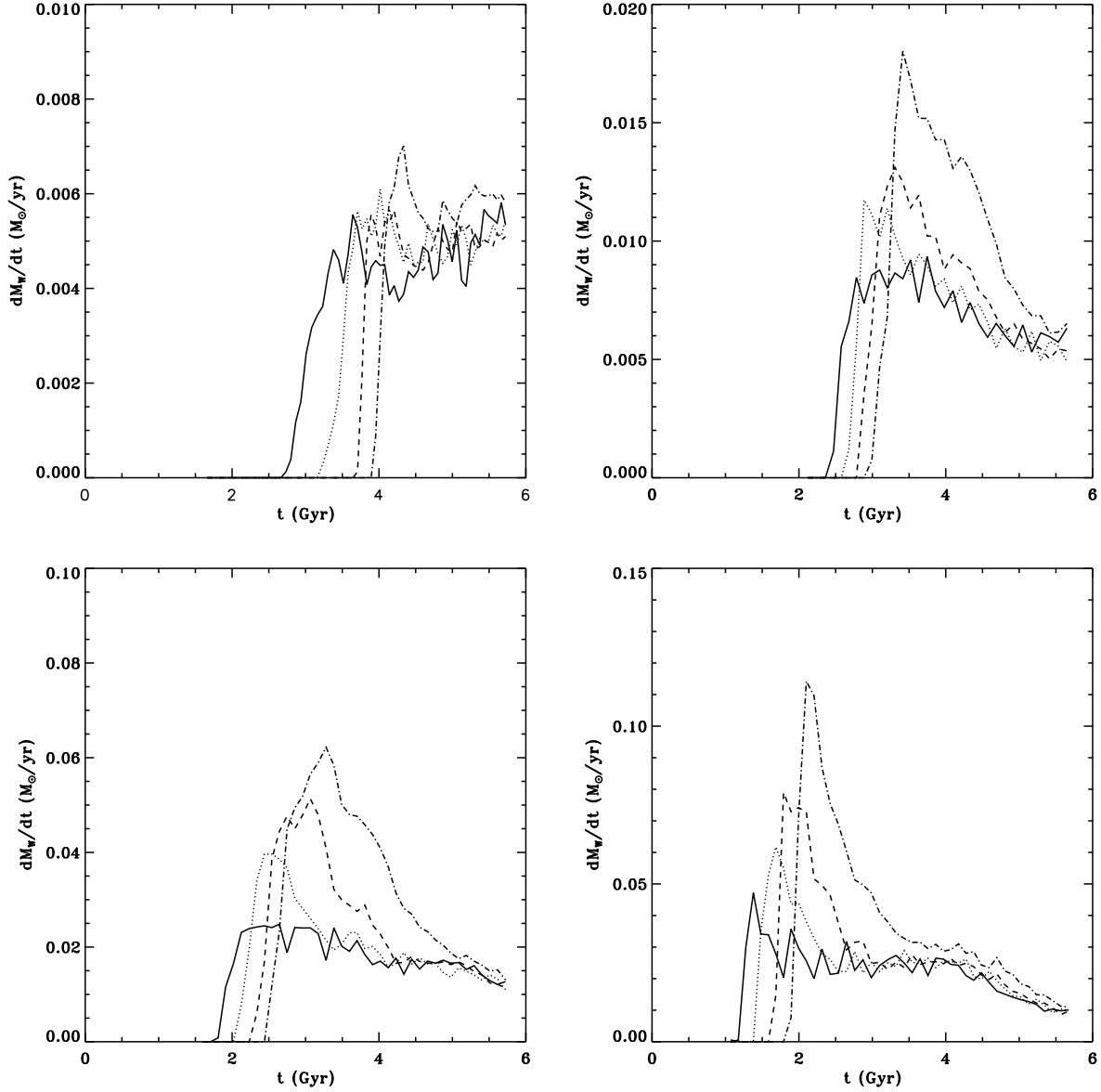


**Fig. 3.** Cut of the gas density (up) and metallicity (bottom) in the Oyz plane for the Sd simulation at different epochs. The two left panels are a zoom 4 times of the simulation box and the two right panels are the entire simulation box. The colour scale gives the density in  $\log(\text{cm}^{-3})$  (up) and the metallicity in  $Z_{\odot}$  (bottom). Note that length scales are not the same in each panel.

is further refined up to 4 or 5 additional levels of refinement ( $\ell_{max} = 11$  or  $12$  corresponding to  $\sim 75$  pc, see table 1 for more details), based on a quasi-Lagrangian strategy, for which a cell is refined if its mass exceeds a threshold of  $2 \times 10^{-6}$  the Virial mass. This refinement strategy verifies the Jeans length criterion inside the galactic disc because of the polytropic component that helps to stabilize against fragmentation. Starting with  $2 \times 10^6$  cells, our simulations usually reach a total of 4 to  $5 \times 10^6$  cells. We end our simulation after a typical integration time of 6 Gyr, which corresponds to the free-fall time of the halo outer boundary at 2 Virial radius.

#### 4. Simulation results

In this section, we present our simulation results for the 2 different halo mass considered in this paper, namely  $10^{10}$  and  $10^{11} M_{\odot}$ . In the first case, a strong wind generally develops after a few Gyr. We will examine in details the wind structure and physical properties. For the second case, because of the ram pressure of infalling gas from the surrounding halo, we see no wind forming: a clear galactic fountain sets in and creates a turbulent, hot atmosphere above (and below) the rotating disc. We will study the structure of this galactic fountain, and



**Fig. 5.** Mean flux of mass outflow calculated for the  $10^{10} M_{\odot}$  halos with  $\lambda = 0.1$  and  $t_0 = 8$  Gyr (top left panel),  $\lambda = 0.1$  and  $t_0 = 3$  Gyr (top right panel),  $\lambda = 0.04$  and  $t_0 = 8$  Gyr (bottom left panel) and  $\lambda = 0.04$  and  $t_0 = 3$  Gyr (bottom right panel), and computed at different radius of the halo:  $r = [5r_s; 7r_s]$  (solid line),  $r = [9r_s; 11r_s]$  (dotted line),  $r = [14r_s; 16r_s]$  (dashed line),  $r = [19r_s; 21r_s]$  (dash-dotted line).

Run	$V_{vir}$ km/s	$\lambda$	$t_0$ Gyr	$\ell_{min}$	$\ell_{max}$	$\Delta x$ pc
Sa	35	0.04	3	7	11	73
Sb	35	0.04	8	7	11	73
Sc	35	0.1	3	7	11	73
Sd	35	0.1	8	7	11	73
La	75	0.04	3	7	12	78
Lb	75	0.04	8	7	12	78
Lc	75	0.1	3	7	12	78
Ld	75	0.1	8	7	12	78

**Table 1.** Parameters used for the simulations performed in this paper.

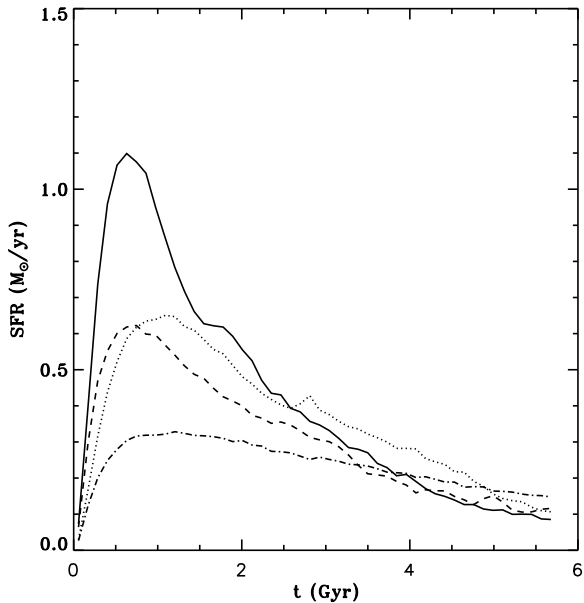
compare the various density and velocity profiles obtained in both cases (wind versus fountain).

#### 4.1. $10^{10} M_{\odot}$ halo

In figure 4, we have plotted the star formation rate (SFR) obtained in the low mass galaxies (Sa, Sb, Sc and Sd, see Table 1 for details). The smallest disc with the highest star formation efficiency (run Sa) shows, as expected, the highest peak in the star formation rate. Using our analytical model, we are in a position to roughly recover the same curve using  $t_* \approx 1$  Gyr.

Run Sb and Sc have roughly the same value for the product  $t_0 \lambda_0$ , and therefore, as predicted by our simple analytical





**Fig. 4.** Star formation rates of the  $10^{10} M_{\odot}$  halos with  $\lambda = 0.04$  and  $t_0 = 3$  Gyr (solid line),  $\lambda = 0.04$  and  $t_0 = 8$  Gyr (dotted line),  $\lambda = 0.1$  and  $t_0 = 3$  Gyr (dashed line) and  $\lambda = 0.1$  and  $t_0 = 8$  Gyr (dash-dotted line).

model (Eq. 6-7), have roughly the same star formation history ( $t_* \simeq 2$  Gyr). The largest disc with the lowest star formation efficiency ( $t_* \simeq 4$  Gyr) has the lowest SFR. At late time, after 3 to 4 Gyr, the 4 curves converge to roughly the same value. This corresponds to the late epoch when star formation is mainly determined by the accretion of gas infalling from the outer halo, as predicted by equation 7.

Figure 2 shows a map of the gas temperature in a cut perpendicular to the galactic plane for run Sd ( $10^{10} M_{\odot}$  with  $\lambda = 0.1$  and  $t_0 = 8$  Gyr) at different epochs. The first panel, at  $t = 1.5$  Gyr, shows no sign of wind, although it corresponds already to the peak of star formation in this galaxy (around  $0.3 M_{\odot}/\text{yr}$ ). We see several hot features trying to break out of the cold disc, but none of them is able to expand significantly, because of the ram pressure of infalling halo material. The second panel corresponds to a later epoch when the star formation rate has dropped by 50%, down to  $0.2 M_{\odot}/\text{yr}$  (see also Figure 3). Nevertheless, since gas accretion has also dropped significantly, two hot and large cavities now emerge out of the disc and propagate in the halo. At  $t = 4.5$  Gyr the galactic wind has finally escaped the halo outer boundaries, and is now in a quasi-permanent regime. This 2 large cavities have created some kind of tunnel, facilitating the ejection of hot, supernovae-driven material from the inner disc, all the way out to the intergalactic medium. At  $t = 6$  Gyr, we see a final snapshot of the galactic wind, with a typical nozzle-like shape (see Fig. 2 and 3). The hot, metal-rich outflowing gas is surrounded by a dense shell of compressed and cooling halo gas. This dense shell will eventually fragment into small clouds that will fall back to the disc. We believe that this dense shell corresponds to the structure seen in many observations around M82 (Heckman et al. 1990;

Shopbell & Bland-Hawthorn 1998; Martin 1998; Hoopes et al. 2003). This also corresponds to the “superwind” geometry described in Tenorio-Tagle & Munoz-Tunon (1998) (see also Veilleux et al. 2005 and references therein).

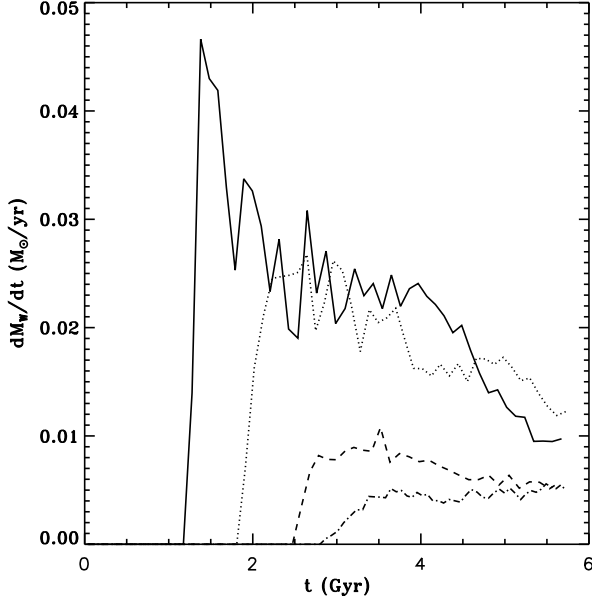
#### 4.2. Wind formation epoch

In order to describe the outflow more quantitatively, we define the net mass flux across a shell of radius  $r_{min}$  and thickness  $\Delta r = r_{max} - r_{min}$  as

$$F = \int_{r_{min}}^{r_{max}} \frac{\rho \mathbf{u} \cdot \mathbf{n}}{\Delta r} 4\pi r^2 dr = F_+ + F_-, \quad (15)$$

where  $\mathbf{n} = \mathbf{r}/r$ . This net flux is further splitted into two different contribution: the positive flux,  $F_+$ , corresponding only to outflowing volume elements, namely those that satisfy  $\mathbf{u} \cdot \mathbf{n} > 0$ , and the negative flux,  $F_-$ , corresponding to inflowing volume elements, satisfying  $\mathbf{u} \cdot \mathbf{n} < 0$ . We focus here on the positive flux, in order to detect the outflow, and to estimate the amount of gas expelled as a function of time. The shell thickness was set to  $\Delta r = 2r_s$ , and we vary the shell radius from  $5r_s$  to  $20r_s$ , in order to define a proper radius where to detect the outflow. Figure 5 shows the positive flux measured at 4 different radii for our 4 low mass galaxies, as a function of time. The wind is detected when the positive flux sharply rise from zero to its maximum value. The smaller the shell radius is, the earlier the wind is detected. For our smallest shell radius,  $r_{min} = 5r_s$ , the measured mass flux remains quasi constant in time, while for larger shell radii, the measured mass flux sharply rises to a much larger value, namely 2 to 3 times the inner flux, then slowly decays to the correct wind mass flux. This means that when the wind breaks out of the halo, a significant fraction (up to  $\sim 2/3$ ) of the wind is filled by the hot gaseous halo and expelled out of the halo boundaries. When the permanent regime sets in at later time, the mass flux is roughly equal to the inner one, independant of the radius of the shell. We therefore consider only the case  $r_{min} = 5r_s$ , as the correct proxy for measuring the mass outflow rate in the wind.

We have plotted the positive flux at  $r_{min} = 5r_s$  for our 4 galaxies in figure 6. The highest SFR galaxy, Sa, corresponds to our earliest wind formation epoch, around  $t_w \simeq 1$  Gyr. According to our analytical model (see figure 1), this corresponds to the hydrodynamical efficiency parameter  $\chi \simeq 0.9\%$ . Although runs Sb and Sc have identical star formation history (see figure 4), they have different wind formation epochs, respectively  $t_w \simeq 2$  and 3 Gyr. This demonstrates that for a given star formation rate, a more compact disc is more efficient in forming galactic winds. According to our analytical recipe, run Sb corresponds to  $\chi \simeq 0.8\%$ , while run Sc barely reaches  $\chi \simeq 0.6\%$ . The final galaxy, Sd, features the latest wind to develop, with  $t_w \simeq 4$  Gyr, corresponding also to a slightly lower hydrodynamical efficiency  $\chi \simeq 0.7\%$ . The main conclusion of this section is that the conversion efficiency of supernovae energy deposition rate into wind luminosity is rather low in this model, around  $\chi \simeq 1\%$ . This efficiency is directly related on the choice of supernovae energy injection rate, a quite uncertain parameter depending on the scale injection energy (here chosen



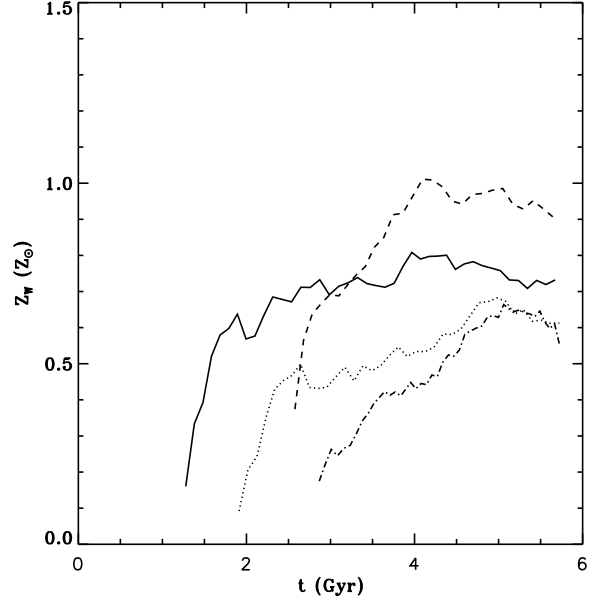
**Fig. 6.** Mean flux of mass outflowing calculated between  $r = [5r_s; 7r_s]$  for the  $10^{10} M_\odot$  halos with  $\lambda = 0.04$  and  $t_0 = 3$  Gyr (solid line),  $\lambda = 0.04$  and  $t_0 = 8$  Gyr (dotted line),  $\lambda = 0.1$  and  $t_0 = 3$  Gyr (dashed line) and  $\lambda = 0.1$  and  $t_0 = 8$  Gyr (dash-dotted line).

equal to 50%). We can see directly from figure 1 that, according to our numerical experiments, no wind will ever form in higher mass halo, as will be demonstrated in the next section.

#### 4.3. Wind efficiency and metallicity

Before describing the simulation results for the  $10^{11} M_\odot$  halo, we would like to characterize the physical properties of the galactic winds we have obtained. It is common to define the wind efficiency as  $\eta_w = F_+/\dot{M}_*$ . In other words, the mass outflow rate is expressed in units of the global star formation rate. From figure 4 and figure 5, we can compute directly  $\eta_w$ : we find the maximum efficiency at late time, when the permanent regime is settled, with values ranging from  $\eta_w \approx 0.05$  for Sc and Sd, up to  $\eta_w \approx 0.1$  for Sa and Sb. Here again, it appears clearly that more compact discs (low spin parameter  $\lambda_0$ ) gives more efficient winds, while extended discs (high spin parameter) are less efficient. All these values are extremely low compared to the wind efficiencies observed by Martin (1999) in Lyman break galaxies with values greater than 1. These high-redshift galaxies are likely massive starbursts, for which our quiescent approach does not apply. On the other hand, we have determined that compact star formation sites result in earlier and stronger winds. Nuclear starbursts can be natural candidates to host very efficient winds with  $\eta_w \geq 1$ , although these bright galaxies are probably rather rare objects in the early universe.

As seen from figure 3, the hot gas carried away in the wind is highly enriched with metals ( $Z_w \approx Z_\odot$ ). We define the average wind metallicity as the positive flux of metals divided by



**Fig. 7.** Mean metallicity of the wind calculated between  $r = [5r_s; 7r_s]$  for the  $10^{10} M_\odot$  halos with  $\lambda = 0.04$  and  $t_0 = 3$  Gyr (solid line),  $\lambda = 0.04$  and  $t_0 = 8$  Gyr (dotted line),  $\lambda = 0.1$  and  $t_0 = 3$  Gyr (dashed line) and  $\lambda = 0.1$  and  $t_0 = 8$  Gyr (dash-dotted line).

the positive flux of mass  $Z_w = (\rho Z u)_+ / F_+$ . Figure 7 shows the mean metallicity of the wind for our various low mass galaxies. We see no clear trend of the wind metallicity as a function of disc size or star formation efficiency. We can only observe that the metallicity of the wind is of about 0.5 to  $1 Z_\odot$ . This rather high value confirms that the wind comes mainly from the gas within the galactic disc that is directly enriched by exploding supernovae.

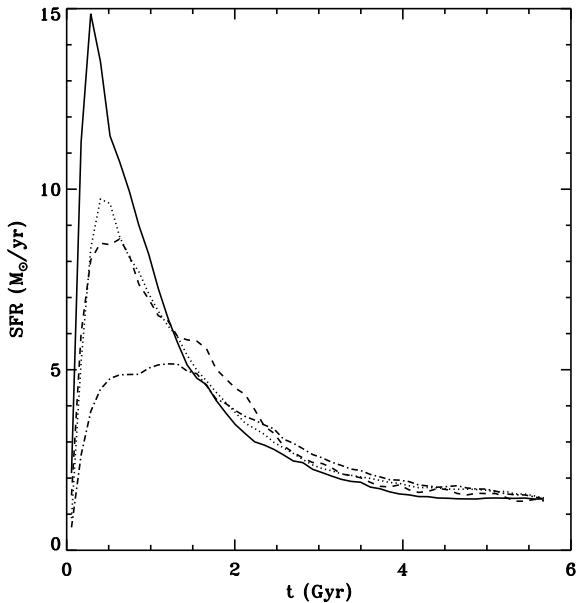
We now briefly discuss how these winds might explain the observed metallicity of the intergalactic medium (IGM). We know from our simulations that the typical mass outflow rate is around  $\dot{M}_w \approx 0.01 M_\odot/\text{yr}$  (see figure 6), with an average wind metallicity of  $Z_w \approx 1 Z_\odot$  and a typical wind velocity around  $u_w \approx 300$  km/s (see figure 13). If we assume that the IGM is photo-ionized to a temperature around  $10^4$ - $10^5$  K, we can compute the volume  $V$  occupied by the wind when the bubble stalls by pressure equilibrium with the IGM as

$$\dot{M}_w u_w^2 t \approx \frac{P_{IGM} V}{\gamma - 1}, \quad (16)$$

assuming the gas is adiabatic. After 6 Gyr, the radius  $R \approx V^{1/3}$  of the expanding wind lies around 1 Mpc. This corresponds to a swept up mass of baryons roughly  $5 \times 10^9 M_\odot$ . If we assume that the metals carried out by the wind perfectly mix with the zero-metallicity IGM, we obtain for the metallicity of the IGM the following very simple relation

$$Z_{IGM} \approx \left( \frac{c_s}{u_w} \right)^2 Z_\odot \approx 10^{-3} - 10^{-2} Z_\odot \quad (17)$$

where  $c_s \approx 10$ -30 km/s is the sound speed of the IGM. Note that this analysis is valid only for the idealized case we have



**Fig. 8.** Star formation rates of the  $10^{11} M_{\odot}$  halos with  $\lambda = 0.04$  and  $t_0 = 3$  Gyr (solid line),  $\lambda = 0.04$  and  $t_0 = 8$  Gyr (dotted line),  $\lambda = 0.1$  and  $t_0 = 3$  Gyr (dashed line) and  $\lambda = 0.1$  and  $t_0 = 8$  Gyr (dash-dotted line).

considered here. Observations of the Ly $\alpha$  forest clouds show that they are metal enriched from  $Z \approx 10^{-3}$  up to  $10^{-2} Z_{\odot}$  (Songaila & Cowie 1996, Ellison et al. 2000). Cosmological simulations also confirm this trend, although the various wind models proposed (including our) might be difficult to reconcile (Gnedin 1998, Thacker et al. 2002).

#### 4.4. $10^{11} M_{\odot}$ halo

We have plotted in figure 8 the star formation history for the 4 galaxies hosted by the  $10^{11} M_{\odot}$  halo. They are qualitatively very similar to the  $10^{10} M_{\odot}$  case, with La and Ld as the most extreme models, and Lb and Lc being quasi undistinguishable. Each galaxy has however a peak star formation rate which is 50% higher than its rescaled  $10^{10} M_{\odot}$  counterpart. The circular velocity in these new galaxies is higher than the low mass ones, while the gas sound speed, because of the polytropic multiphase model, remains roughly constant, around 10 km/s, so that the disc is now much thinner. This results in higher total star formation efficiencies. Using our analytical model, we can reproduce the 3 star formation histories using for run Sa, Sb (and Sc) and Sd the values  $t_* \approx 0.75, 1.5$  and 3 Gyr respectively.

The other striking differences with low mass galaxies is the clear presence of an accretion shock surrounding the disc, and the notable absence of galactic winds. We observe instead a nice galactic fountain, with plumes of hot gas rising above the disc, cooling down and falling back to the disc (see figure 9).

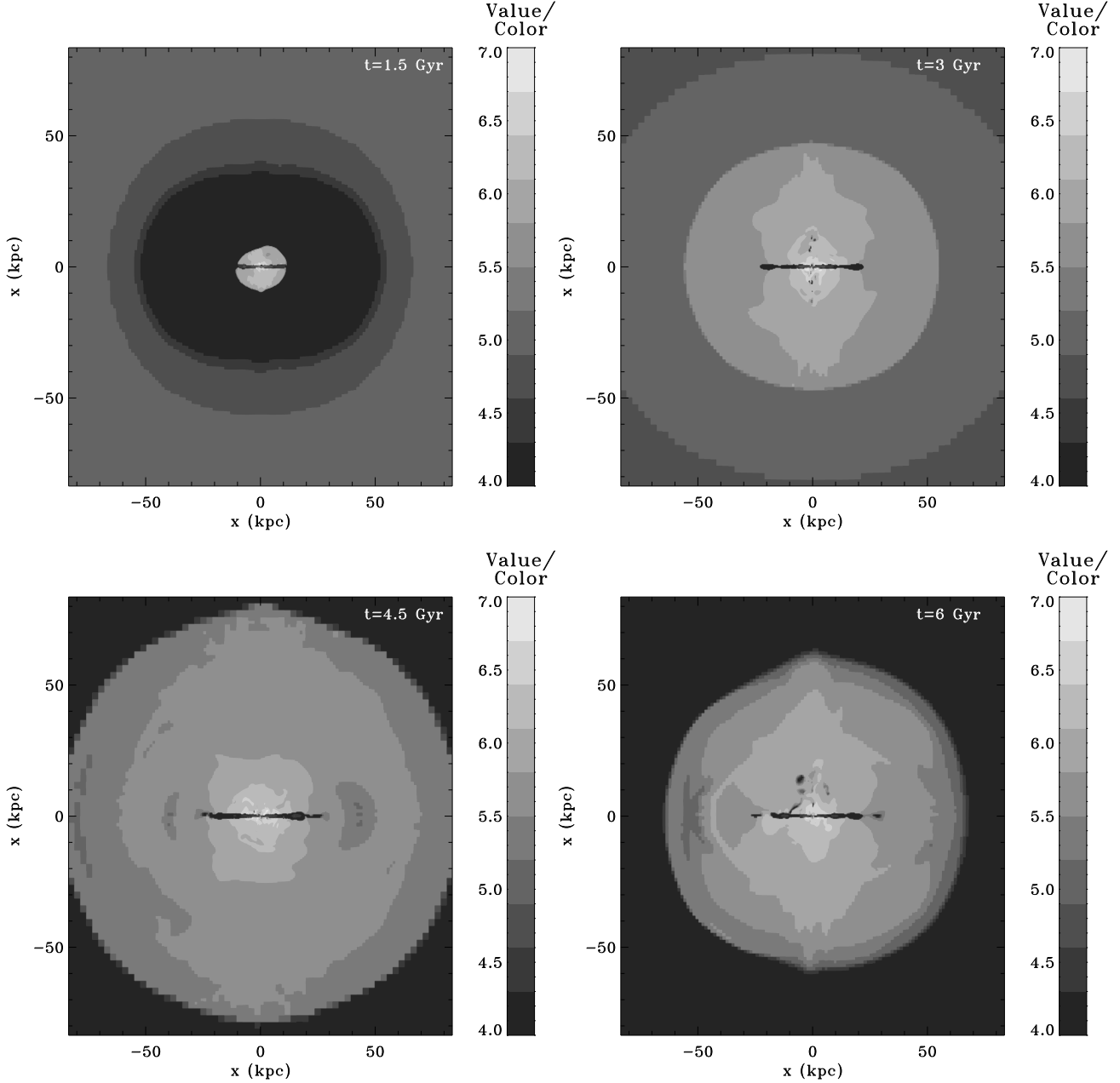
The formation of an accretion shock above a certain halo mass is now well understood: the stability analysis of a radiative shock wave, performed by Birnboim & Dekel (2003), ex-

plains why for low mass halo no accretion shock forms. The mass threshold above which the shock forms depends on the metallicity of the cooling gas. We see from figure 9 that in our simulation the shock heated gas above and below the disc is metal poor: in this case, Birnboim & Dekel (2003) found a critical mass for shock appearance slightly below  $10^{11} M_{\odot}$  for a zero metallicity gas, in good agreement with our numerical experiment. The accretion shock converts the infall kinetic energy into internal energy, so that our analytical treatment of wind formation remains valid, replacing the ram pressure by the thermal pressure. The accretion shocks reaches a maximum radius of  $\sim 75$  kpc what is  $\sim 2/3 r_{vir}$  quite the same virial radius fraction than the one found in Birnboim & Dekel (2003) for the cooling case after  $\sim 4 - 5$  Gyr, and finally contracts down to  $\sim 50$  kpc at later time. Although the disc is strongly perturbed by supernovae blast waves, no galactic winds can form, in agreement with our simple analytical prediction, if we assume for the hydrodynamical efficiency  $\chi \approx 1\%$ , a value close to the one obtained in the  $10^{10} M_{\odot}$  case. The atmosphere above and below the disc is however highly turbulent and perturbed by buoyantly rising plumes. This convective flow is more visible in the metallicity map of figure 10.

Metals are confined within a rather small distance relative to the disc plane, namely  $\sim 20$  kpc. In this case, the thermal pressure of the shock-heated halo gas is very efficient in confining the galactic wind. Figure 11 compares the metallicity profiles of the gas between the  $10^{10} M_{\odot}$  halo and the  $10^{11} M_{\odot}$  halo (run Sd versus run Ld) at different epochs. We define for that purpose a mass-weighted average metallicity as a function of height and a mass-weighted metallicity as a function of radius as

$$Z(z) = \frac{\int_0^{r_d} \rho Z 2\pi r dr}{\int_0^{r_d} \rho 2\pi r dr} \quad \text{and} \quad Z(r) = \frac{\int_{-h_d}^{+h_d} \rho Z 2\pi dz}{\int_{-h_d}^{+h_d} \rho 2\pi dz} \quad (18)$$

where  $\rho$  is the gas density and  $Z$  is the gas metallicity. The disc radius  $r_d$  is taken equal to  $r_s$  and the disc scale height  $h_d = r_d/10$ . In the wind case, metals are entrained at very large distance above and below the disc in the wind nozzle. The metallicity is roughly solar and tends to saturate at that value. In the fountain case, the metallicity in the halo remains rather low, roughly one tenth solar. On the other hand, the gas metallicity in the central part of the disc is significantly higher than in the wind case ( $Z_{gas} \approx 1.2 Z_{\odot}$  versus  $Z_{gas} \approx 0.6 Z_{\odot}$ , see figure 11). This could be interpreted as the effect of the wind removing metals preferentially from the bulge of the galaxy. In the outer parts of the disc, however, the gas metallicity is very similar in both cases. This could also be interpreted as a higher star formation efficiency in the high mass case, especially in the central parts where the density is higher. As explained in great details in Dalcanton (2006), it is quite difficult to disentangle the influence of star formation, gas infall and winds on metal enrichment in galaxies. We believe that, even in our idealized case, the metal content in our galaxies is determined by a subtle balance between these 3 processes.

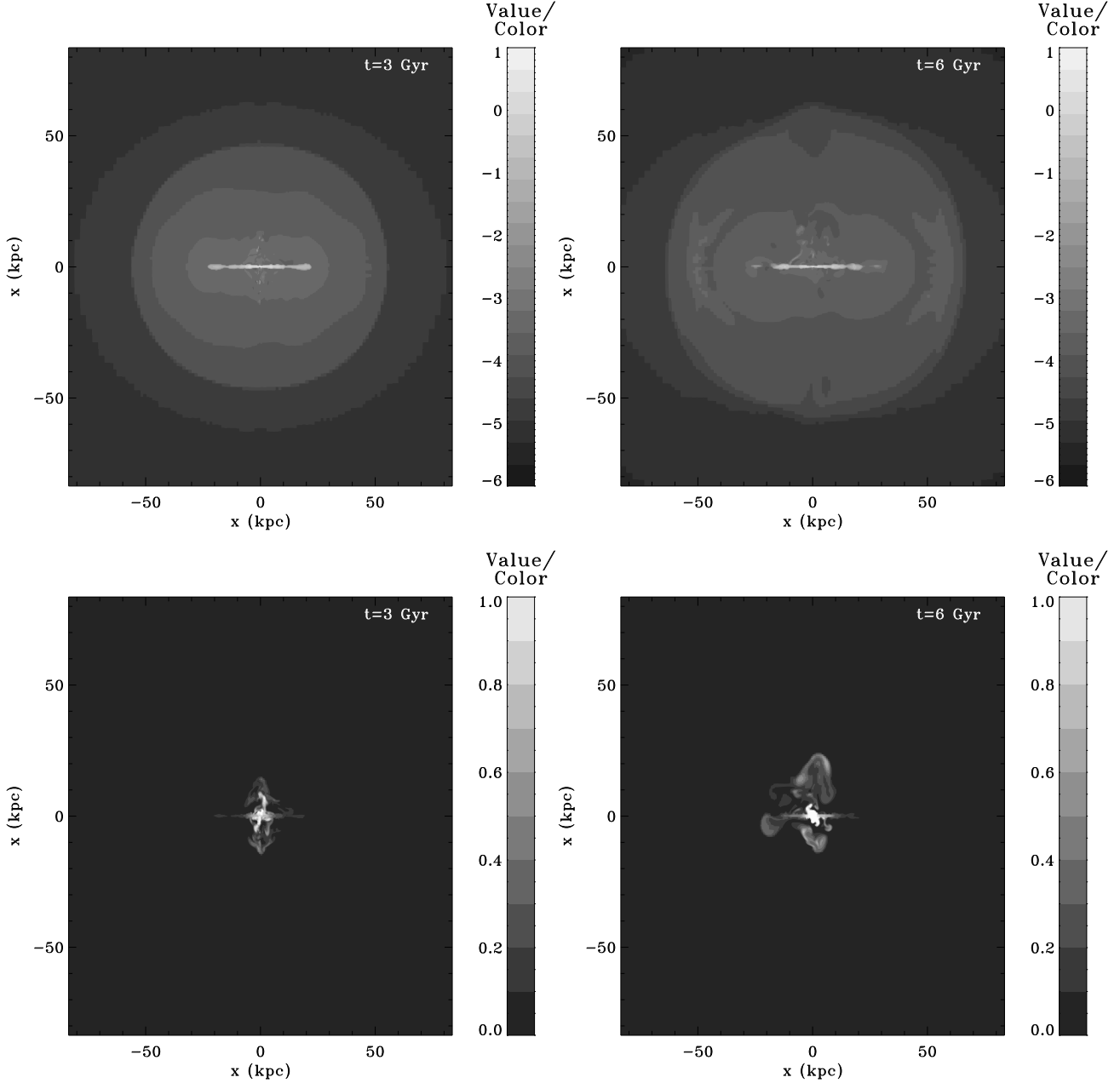


**Fig. 9.** Cut of the gas temperature in the Oyz plane for the Ld simulation at different epochs. The 4 pannels are a zoom 4 times of the simulation box. The colour scale gives the temperature in  $\log(K)$ . Note that length scales are not the same in each pannel.

#### 4.5. Stellar distribution and metallicity

Figure 12 shows the surface density profile of stars and the corresponding mass-weighted average metallicity as a function of the disc radius. In both cases, we see a clear exponential disc, with a bulge-like feature in the central part and a sharp decline of surface density in the outer part. We find a bulge radius of roughly 1 kpc and a disc exponential radius of 2.5 kpc for the low mass system and of 4 kpc for the high mass one. We note that the outer radius of the low mass case (8 kpc) is smaller than the rescaled outer radius of its high-mass counterpart (10 kpc). This is due to the different disc thickness, so that for a given surface density, the threshold for star forma-

tion is reached more easily by the most massive galaxy. This confirms the fact that the polytropic multiphase model has a great importance in regulating the gaseous disc properties and therefore star formation efficiencies. The break in the stellar surface density corresponds roughly to the  $\bar{\rho}_0$  average radius where the gas reaches the density threshold of star formation ( $\rho_0 = 0.1 \text{ H.cm}^{-3}$ ). Note that the overall amount of stars is very similar in both cases, demonstrating that the presence of the galactic wind has no impact on the amount of baryons locked into long-lived stars. This is consistent with the rather low wind efficiencies we have measured ( $\eta_w \approx 10\%$ ). Our simulations will not solve the overcooling problem. More importantly, the stellar metallicity seems also to be rather insensitive



**Fig. 10.** Cut of the gas density (up) and metallicity (bottom) in the Oyz plane for the Ld simulation at different epochs. The two left panels are a zoom 4 times of the simulation box and the two right panels are the entire simulation box. The colour scale gives the density in  $\log(\text{cm}^{-3})$  (up) and the metallicity in  $Z_{\odot}$  (bottom). Note that length scales are not the same in each panel.

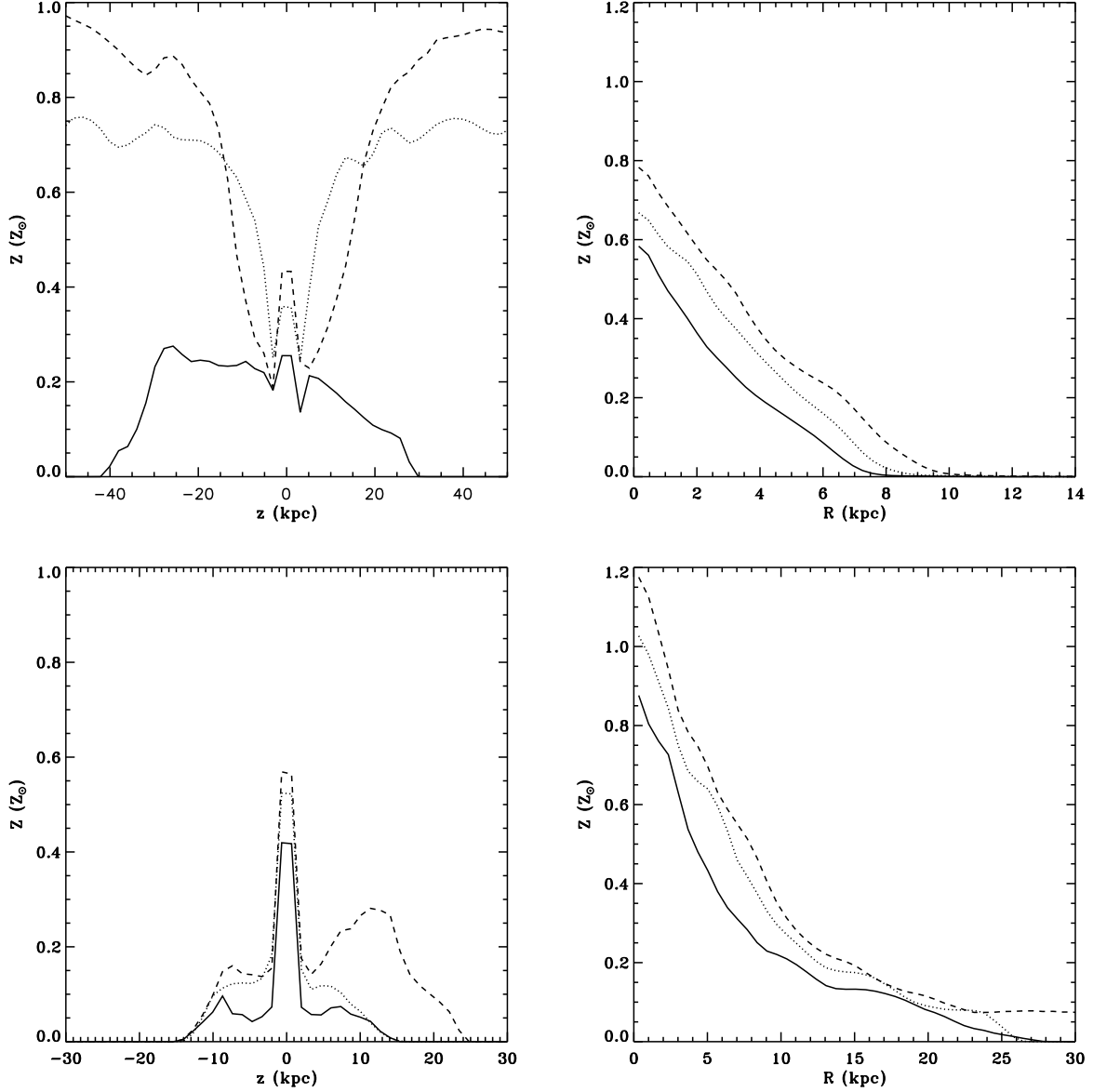
to the “wind versus fountain” scenario. We have  $Z_* \simeq 0.4 Z_{\odot}$  in the central part of our  $10^{11} M_{\odot}$  halo galaxy, while we have  $Z_* \simeq 0.3 Z_{\odot}$  in the low mass case. Note that we have computed a mass-weighted stellar metallicity so we are more sensitive to the overall star formation history rather than to the latest OB stars formed. Nevertheless, our wind model has no dramatic impact on the simulated stellar population, although there is a weak trend for low mass galaxies to have slightly smaller stellar metallicities than high mass. Here again, as in Dalcanton (2006), it is quite difficult to conclude whether this trend is due to a wind or to a smaller star formation efficiency.

#### 4.6. Gas kinematics

The most striking and observable differences between the wind and the fountain can be found in the gas kinematics in the halo atmosphere, below and above the disc. In order to mimic as much as possible the observational signatures of our simulated galaxies, as seen by a spectral lines analysis from a distant observer whose line of sight is aligned with the disc rotation axis, we define the average metal velocity as

$$v_{\perp}(z) = \frac{\int_0^{r_d} \rho Z v_z 2\pi r dr}{\int_0^{r_d} \rho Z 2\pi r dr} \quad \text{and} \quad v_{\perp}^2(z) = \frac{\int_0^{r_d} \rho Z v_z^2 2\pi r dr}{\int_0^{r_d} \rho Z 2\pi r dr} \quad (19)$$





**Fig. 11.** Mean metallicity of the gas as a function of the projected height (left) and radius (right) for Sd simulation (up) and the Ld simulation (bottom) at different times:  $t = 3$  Gyr (solid line),  $t = 4.5$  Gyr (dotted line),  $t = 6$  Gyr (dashed line).

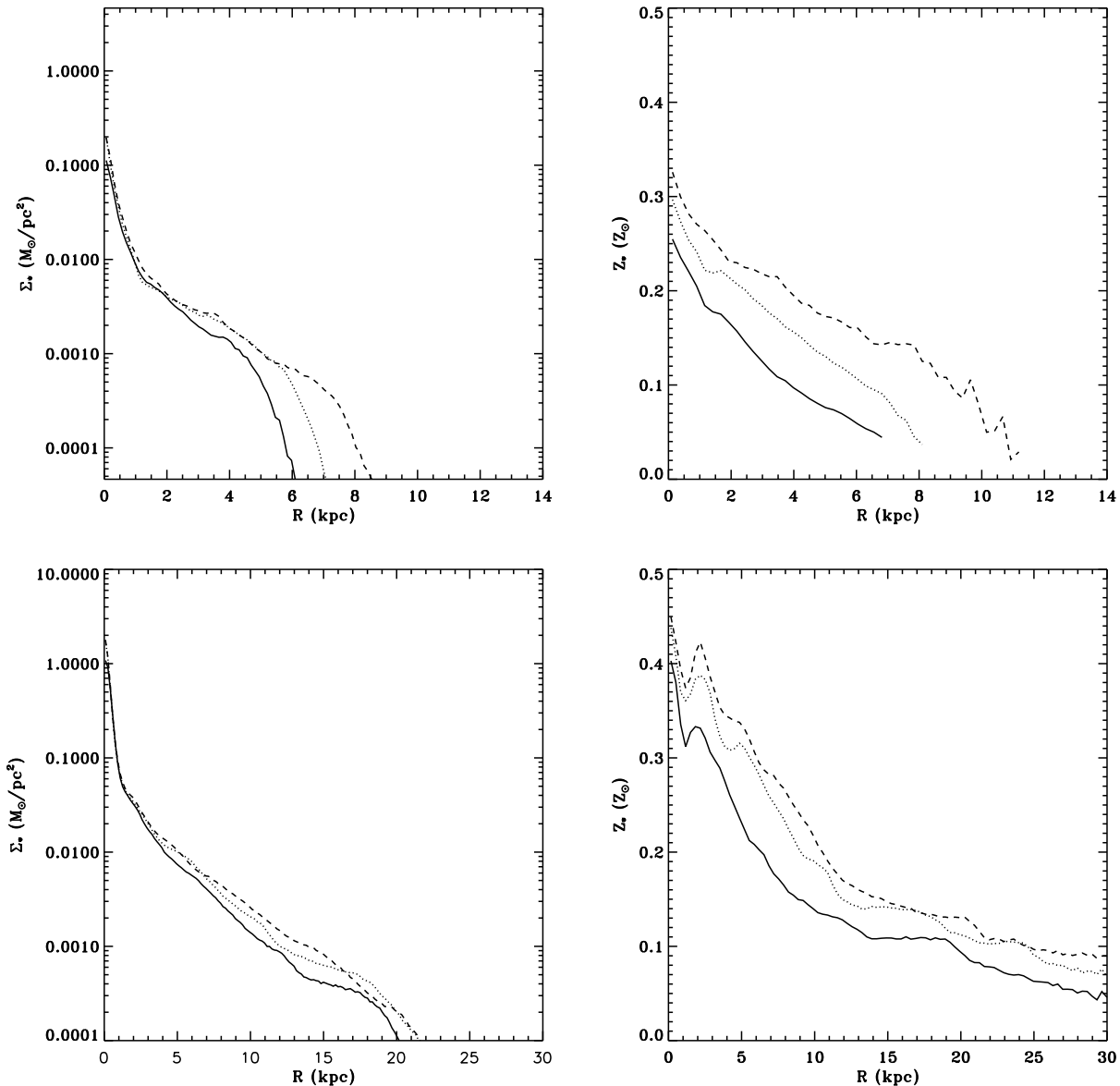
The velocity dispersion and the sound speed are defined as usual by

$$\sigma_{\perp}^2(z) = v_{\perp}^2(z) - v_{\perp}(z)^2 \quad \text{and} \quad c_s^2(z) = \frac{\int_0^{r_d} \rho Z c_s^2 2\pi r dr}{\int_0^{r_d} \rho Z 2\pi r dr} \quad (20)$$

Figure 13 shows these various profiles as a function of height for the wind case (left) and the fountain case (right). In the former case, we clearly see a strong symmetrical expansion of the metals, with a strong acceleration up to 200 km/s in 10 kpc, then a plateau afterwards when the wind reaches its terminal velocity. In the latter case, on the contrary, we see a converging flow, with a velocity of the order of 100 km/s, similar to the velocity of infalling halo material. At an altitude of 20 kpc, this bulk infall velocity declines and in the same time, the velocity dispersion rises: we are entering the galactic fountain for which

the flow is subsonic and highly turbulent. Note that in the wind case, the velocity dispersion also rises up to 100 km/s as we approach the disc. This can't be attributed to a turbulent flow, since it would have been supersonic (the sound speed is around 50 km/s). It is in fact due to a geometric effect: as soon as the wind expansion velocity remains subsonic, below an altitude of 10 kpc, the section of the wind increases with height. When the wind turns supersonic, around  $z \simeq 10$  kpc, the section of the wind is forced to remain constant. This process is responsible for the typical nozzle-like shape of the wind. At low altitude, since the wind is expanding quasi isotropically, this mimics a supersonic velocity dispersion that vanishes at higher altitude when the wind flow is quasi parallel to the galaxy rotation axis.

The  $10^{10} M_{\odot}$  halo is characterized by a fast supersonic flow, expanding quasi-isotropically at low altitude but quasi-parallel



**Fig. 12.** Surface density (left) and mean metallicity (right) of the stars as a function of the projected radius for the Sd simulation (up) and the Ld simulation (bottom) at different times:  $t = 3$  Gyr (solid line),  $t = 4.5$  Gyr (dotted line),  $t = 6$  Gyr (dashed line).

to the rotation axis at higher altitude. The  $10^{11} M_\odot$  halo, on the contrary, is characterized by a converging pattern quickly thermalized close to the disc in a subsonic turbulent flow.

We applied the method used by Prochaska et al. (2007) to test whether our simulations show velocity dispersion in the metal absorption lines of QSOs wider than the velocity dispersion expected from pure gravitational processes (of the order of the galaxy circular velocity). In this way, we can test whether our winds would have been detected. In their paper, Prochaska et al. (2007) use  $\Delta v_{90}$  as the velocity interval encompassing 90% of the mass of a given metal absorber. We simulate an observer watching the galaxy along the direction of the wind and found that 90% of the mass of metals corresponds to metals with velocity lower than  $5 \text{ km.s}^{-1}$ . This is in agreement with our findings that wind efficiencies are very low in our model.

## 5. Conclusion

In comparison with Springel & Hernquist (2003), we have very similar initial conditions but a lower wind for  $\eta_w$ . We have therefore more difficulties to solve the 'over-cooling' problem. The explanation has two reasons : on the one hand Springel & Hernquist (2003) model supernovae explosions with a phenomenological approach so that every single ejecta in the disc entirely participate to the large scale out-flow. In this work, our goal was to model supernovae explosions with a self-consistent Sedov blast approach where ejectas are tightly coupled to the galactic disc (injecting energy in the galactic turbulent cascade). On the second hand, we have a cooling function depending on metallicity: this increases radiative losses significantly compared to Springel & Hernquist (2003). Kobayashi et al. (2006) performed SPH simulations of

isolated disks with star formation and self-consistent supernovae feedback, metal dependent cooling and the same set of initial conditions than ours. They observed very high mass ejection rates for the  $10^{10} M_{\odot}$  halo with a total mass ejected rising to 80% of the total baryons mass. They noted however a strong dependance of their results to numerical resolution and, more importantly, to the size of the feedback radius (it can be as high as 10 kpc !). They even observed a faint wind in the  $10^{12} M_{\odot}$  halo case. Tasker & Bryan (2006) have simulated an pre-formed isolated disk with the AMR code ENZO with star formation and self-consistent supernovae feedback released in pure thermal form over one dynamical time scale. They succeed in shutting down star formation but they didn't take into account infalling gas from the halo. Fujita et al. (2004) constructed an analytical prescription to constrain the mass of the halos that can form large-scale outflow with starburst physics. They have performed several sets of 2D grid simulations of a pre-formed galaxy with an imposed analytical infall model and found results very similar to ours (see Figure 1 in Fujita et al. 2004) : the key role played by infall and a very low hydrodynamical efficiency for quiescent modes of star formation.

Using a quiescent model of star formation in isolated galaxies, self-consistently simulated from a cooling NFW halo, we have studied the conditions for a galactic wind to break-out of the ram-pressure exerted by the infalling gas. Our simulations have shown that a galactic super wind forms in halos of  $10^{10} M_{\odot}$  and no wind can form in halo of mass greater than  $10^{11} M_{\odot}$ , even for our most favorable couple of halo parameters ( $\lambda = 0.04$ ,  $t_0 = 3$  Gyr). For large galaxies, a galactic fountain appears and expels metal rich gas that cannot escape from the galaxy. Those results are consistent with our toy model that predicts no wind in those galaxies if the hydrodynamical efficiency is of a few percents as suggested by our simulations. Using this rather simple toy model, we understand this failure as due to the ram pressure of infalling material confining the outflowing wind.

Although galactic winds develop in  $10^{10} M_{\odot}$  halos, they are not sufficient to explain the "overcooling" problem. The mass ejection efficiency obtained in our simulations  $\sim 3 - 12\%$  is one order of magnitude lower than that observed by Martin (1999) in starburst galaxies. The main conclusion of this paper is that with a self-consistent treatment of supernovae feedback in idealized simulations of quiescent star-forming galaxies we cannot reproduce the observed mass ejection rates and therefore reduce the cold baryon fraction. On the contrary, supernovae feedback is an efficient process for metal enrichment of the IGM. Cosmological simulations with supernovae feedback could test this scenario in a more quantitative way. Using a more realistic cosmological setting may result into non-spherical accretion flows, and therefore to a less stringent criterion for a wind to break-out.

Considering that small mass galaxies preferentially forms along filaments, non-spherical accretion should allow the wind to break-out more easily and probably to be more efficient. For more massive galaxies, however, the accretion of gas into the disc is likely to proceed through a Virial shock, and not along filaments (Dekel & Birnboim 2006). Our conclusions should therefore remain valid, so that massive galaxies should

not provide winds powerfull enough to enrich the IGM by SNe feedback (but other way of enrichment like gas stripping from satellites can be considered like in Schindler et al. 2005). Dekel & Birnboim (2006) have computed the critical mass between filamentary accretion and Virial shocks in the range  $10^{11} - 10^{12} M_{\odot}$  depending on the metallicity of the infalling gas. Combining their results with ours, we can conclude that halos larger than a few  $10^{11} M_{\odot}$  cannot produce galactic winds, even within a realistic cosmological environment. At higher redshift, accretion rates are believed to be higher than the one we have considered here, so that galactic winds at earlier epochs are even more unlikely. A proper modelling of starburst (yet to be invented) might also provide an easier route for increasing the feedback efficiency of supernovae-driven outflows.

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## Appendix A: Numerical scheme for Supernovae

We define the mass vanished in the star formation process as :

$$(\Delta m_g)_{SF} = m_*(1 + \eta_{SN} + \eta_W), \quad (\text{A.1})$$

where the second term of the right part is equal to the total debris mass  $m_d$ . We have also introduce the mass loading factor  $\eta_W$ , the parameter that determines the gas mass carried in debris. And  $m_*$  is equal to

$$m_* = \frac{\rho_0 \delta x^3}{1 + \eta_{SN} + \eta_W}, \quad (\text{A.2})$$

$\delta x^3$  is the volume of the cell where the star particle is spawned. Then if we introduce the fraction of debris particles :

$$f_d = \frac{\eta_{SN} + \eta_W}{1 + \eta_{SN} + \eta_W}, \quad (\text{A.3})$$

we can simplify equation (A.1) by saying that a fraction  $f_d$  of the gas consumed goes into debris particles, and a fraction  $1 - f_d$  goes into star particle.

The maximum speed of a debris is given by

$$u_d = \frac{u_{SN}}{\sqrt{1 + \eta_W/\eta_{SN}}}. \quad (\text{A.4})$$

$u_{SN}$  is the typical velocity corresponding to the kinetic energy released in one single supernova explosion ( $u_{SN} \simeq 3200 \text{ km.s}^{-1}$ ).

Thus, the energy released to the gas by the debris is

$$E_d = \eta_{SN} \frac{m_*}{M_{SN}} E_{SN}. \quad (\text{A.5})$$

$M_{SN}$  and  $E_{SN}$  are respectively the typical progenitor mass and energy of an exploding type II supernova (i. e.  $M_{SN} = 10 M_{\odot}$  and  $E_{SN} = 10^{51} \text{ erg}$ ). The energy is independent from the mass loading factor. Then, we have to compute the moment conservation for constraining this free parameter. We want to reproduce a Sedov blast wave, and so the velocity of the gas when

the debris are coupled with it is the Sedov speed of the shock propagation.

$$u_{Sedov} = \beta \frac{2}{5} \left( \frac{E_d}{\rho_0 \Delta x^3} \right)^{1/2}, \quad (\text{A.6})$$

$\beta$  is a multiplication factor nearly equal 1, we admit it is equal to 1,  $\rho_0$  is the density of the gas where the explosion takes place and  $\Delta x$  is the radius of the shock from the center of the explosion. We can also express  $u_{Sedov}$  in terms of  $u_{SN}$ :

$$u_{Sedov} = \frac{\sqrt{2}}{5} \left[ \eta_{SN} \left( \frac{\delta x}{\Delta x} \right)^3 \frac{1}{1 + \eta_{SN} + \eta_W} \right]^{1/2} u_{SN} \quad (\text{A.7})$$

The Sedov blast wave is only valid in an homogeneous fluid, then it seems surprising to apply it in regions of star formation where the gas condenses in high overdensity regions of molecular clouds. That is the reality, but our simulations do not reproduce such high resolution features, and the assessment on the polytropic trend of the gas at high densities implies more pressurized regions in the disc. This behavior at high densities transform high over-density regions of star formation into smoother regions at the scale of one cell of the numerical grid. That is the reason that allows us to use the Sedov blast wave model. We want the gas to have the moment of a Sedov blast wave, so the conservation of the moment gives us :

$$m_d u_d = (m_g + m_d) \bar{u}_{Sedov}, \quad (\text{A.8})$$

where  $m_g = 4/3 \pi \rho_0 \Delta x^3$  is the gas mass where the supernovae explosion occur, and  $\bar{u}_{Sedov} \simeq u_{Sedov}/3$  is the mean velocity of the gas carried along the shock.

As we know all the quantities of this last equation, we can constraint our free parameter  $f_d$ , and obtain

$$\left[ 1 + \frac{1}{f_d} \frac{4}{3} \pi \left( \frac{\Delta x}{\delta x} \right)^3 \right] \frac{\sqrt{2}}{15} \left[ f_d \left( \frac{\delta x}{\Delta x} \right)^3 \right]^{1/2} = 1. \quad (\text{A.9})$$

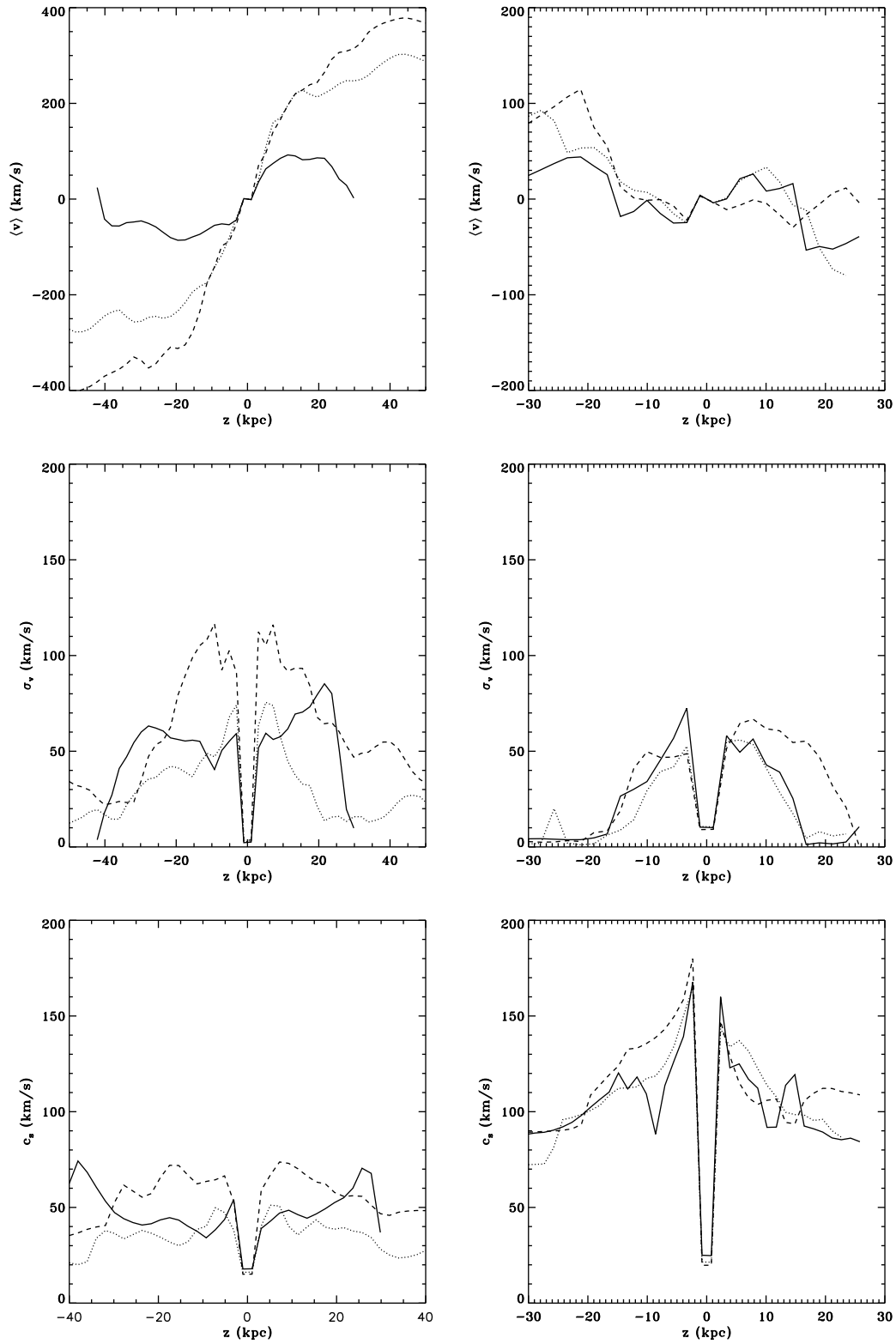
If we want that our debris propagates to the nearest cells, let us assume that  $\Delta x \simeq 1.5$ , and we find the optimal value for  $f_d \simeq 0.5$ . Thus for a typical value of  $\eta_{SN} = 0.1$  we find for the mass loading factor, our only free parameter,  $\eta_W \simeq 1$ .

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**Fig. 13.** Mean velocity (up), velocity dispersion (middle), sound speed (bottom) of the gas as a function of the projected height for the Sd simulation (left) and the Ld simulation (right) at different times:  $t = 3$  Gyr (solid line),  $t = 4.5$  Gyr (dotted line),  $t = 6$  Gyr (dashed line).