The Astrophysical Journal Letters, 715:L6–L11, 2010 May 20 doi:10.1088/2041-8205/715/1/L6

A $z = 1.82$ ANALOG OF LOCAL ULTRA-MASSIVE ELLIPTICAL GALAXIES*

M. Onodera[1], E. Daddi[1], R. Gobat[1], M. Cappellari[2], N. Arimoto[3,4], A. Renzini[5], Y. Yamada[3], H. J. McCracken[6], C. Mancini[5], P. Capak[7], M. Carollo[8], A. Cimatti[9], M. Giavalisco[10], O. Ilbert[11], X. Kong[12], S. Lilly[8], K. Motohara[13], K. Ohta[14], D. B. Sanders[15], N. Scoville[16], N. Tamura[17], and Y. Taniguchi[18]

[1] CEA, Laboratoire AIM-CNRS-Université Paris Diderot, Irfu/SAp, Orme des Merisiers, F-91191 Gif-sur-Yvette, France; masato.onodera@cea.fr
[2] Sub-department of Astrophysics, University of Oxford, Denys Wilkinson Building, Keble Road, Oxford OX1 3RH, UK
[3] National Astronomical Observatory of Japan, Osawa 2-21-1, Mitaka, Tokyo, Japan
[4] Graduate University for Advanced Studies, Osawa 2-21-1, Mitaka, Tokyo, Japan
[5] INAF-Osservatorio Astronomico di Padova, Vicolo dell'Osservatorio 5, I-35122 Padova, Italy
[6] Institut d'Astrophysique de Paris, UMR7095, Université Pierre et Marie Curie, 98 bis Boulevard Arago, 75014 Paris, France
[7] Spitzer Science Center, California Institute of Technology 220-06, Pasadena, CA 91125, USA
[8] Institute for Astronomy, ETH Zurich, Wolfgang-Pauli-strasse 27, 8093 Zurich, Switzerland
[9] Dipartimento di Astronomia, Università di Bologna, Via Ranzani 1, 40127 Bologna, Italy
[10] Department of Astronomy, University of Massachusetts, Amherst, MA, USA
[11] Laboratoire d'Astrophysique de Marseille, BP 8, Traverse du Siphon, 13376 Marseille Cedex 12, France
[12] Center for Astrophysics, University of Science and Technology of China, Hefei, Anhui 230026, China
[13] Institute of Astronomy, University of Tokyo, Mitaka, Tokyo, Japan
[14] Department of Astronomy, Kyoto University, Kyoto 606-8502, Japan
[15] Institute for Astronomy, University of Hawaii, 2680 Woodlawn Drive, Honolulu, HI 96822, USA
[16] California Institute of Technology, MC 105-24, 1200 East California Boulevard, Pasadena, CA 91125, USA
[17] Subaru Telescope, National Astronomical Observatory of Japan, 650 North A'ohoku Place, Hilo, HI 96720, USA
[18] Research Center for Space and Cosmic Evolution, Ehime University, 2-5 Bunkyo-cho, Matsuyama 790-8577, Japan

Received 2009 December 3; accepted 2010 April 12; published 2010 April 26

ABSTRACT

We present observations of a very massive galaxy at $z = 1.82$ that show that its morphology, size, velocity dispersion, and stellar population properties are fully consistent with those expected for passively evolving progenitors of today's giant ellipticals. These findings are based on a deep optical rest-frame spectrum obtained with the Multi-Object InfraRed Camera and Spectrograph on the Subaru Telescope of a high-z passive galaxy candidate (pBzK) from the COSMOS field, for which we accurately measure its redshift of $z = 1.8230$ and obtain an upper limit on its velocity dispersion $\sigma_\star < 326$ km s^{-1}. By detailed stellar population modeling of both the galaxy broadband spectral energy distribution and the rest-frame optical spectrum, we derive a star formation-weighted age and formation redshift of $t_{\rm sf} \simeq 1$–2 Gyr and $z_{\rm form} \simeq 2.5$–4, and a stellar mass of $M_\star \simeq (3\text{–}4) \times 10^{11}\ M_\odot$. This is in agreement with a virial mass limit of $M_{\rm vir} < 7 \times 10^{11}\ M_\odot$, derived from the measured $\sigma_\star$ value and stellar half-light radius, as well as with the dynamical mass limit based on the Jeans equations. In contrast to previously reported super-dense passive galaxies at $z \sim 2$, the present galaxy at $z = 1.82$ appears to have both size and velocity dispersion similar to early-type galaxies in the local universe with similar stellar mass. This suggests that $z \sim 2$ massive and passive galaxies may exhibit a wide range of properties, then possibly following quite different evolutionary histories from $z \sim 2$ to $z = 0$.

Key words: galaxies: evolution – galaxies: formation – galaxies: high-redshift

Online-only material: color figures

1. INTRODUCTION

Understanding the formation of massive elliptical galaxies remains a crucial unsolved issue of galaxy evolution. The recent discovery and the first redshift measurements, through deep ultraviolet (UV) rest-frame spectroscopy, of a substantial population of passively evolving galaxies at $z > 1.4$ (e.g., Cimatti et al. 2004; McCarthy et al. 2004; Daddi et al. 2005) have shown that quenching of star formation (SF) in the most massive galaxies was already well under way at $z \simeq 2$.

A puzzling property of such objects has been revealed soon afterward with some of them being found to have a factor of $\simeq$2–5 smaller effective radii compared to local early-type galaxies (ETGs) of the same stellar mass (e.g., Daddi et al. 2005; Trujillo et al. 2006; Longhetti et al. 2007; Cimatti et al. 2008; van Dokkum et al. 2008), implying that they are $\gtrsim$10 times denser than their possible descendants in the local universe. Several alternative mechanisms have been proposed to make such compact ETGs grow in size so as to finally meet the properties of local ETGs (e.g., Khochfar & Silk 2006; Fan et al. 2008; Naab et al. 2009; La Barbera & de Carvalho 2009; Feldmann et al. 2010), but no general consensus has yet emerged.

On the other hand, ETGs at $z > 1.4$ with large effective radii, comparable to the local ETGs, have also been found (e.g., Mancini et al. 2010; Saracco et al. 2009, see also Daddi et al. 2005), indicating a diversity of structural properties in the ETG population at $z \simeq 2$. Moreover, possible effects have also been discussed that could bias size estimates toward lower values (e.g., Daddi et al. 2005; Hopkins et al. 2010; Mancini et al. 2010; Pannella et al. 2009).

An independent way to check these issues is by measuring stellar velocity dispersions ($\sigma_\star$): if high-z ETGs are really super-dense, their $\sigma_\star$ should be much higher than that of local ETGs of the same mass. Cenarro & Trujillo (2009) and Cappellari et al. (2009) measured $\sigma_\star$ from deep UV rest-frame spectroscopy of a sample of ETGs at $1.4 < z < 2.0$ from the GMASS survey

* Based on data collected at the Subaru Telescope, which is operated by the National Astronomical Observatory of Japan (S09A-043).

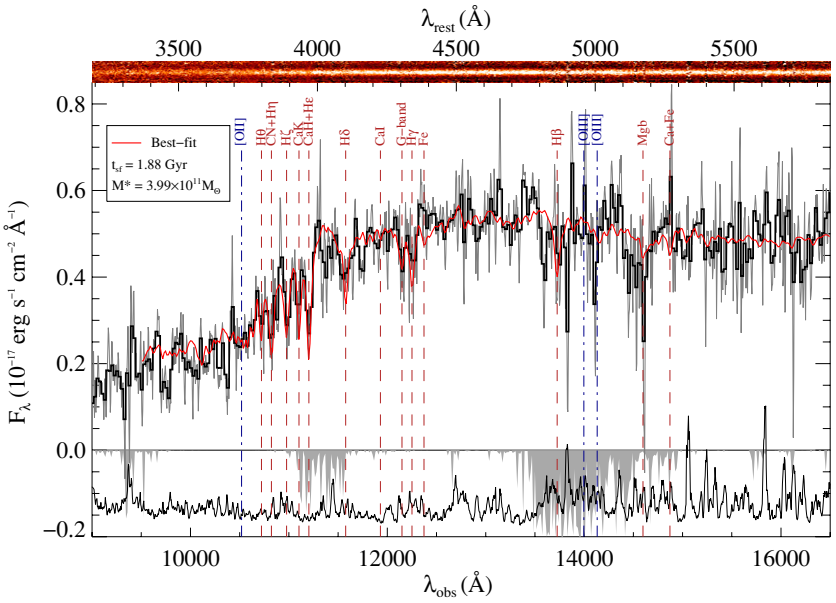

Figure 1. MOIRCS spectrum of 254025 at $z_{\rm spec} = 1.82$. Top: $4''$ two-dimensional spectrum; middle: one-dimensional spectrum without smoothing (gray line) and with a 25 Å binning (black thick line); bottom: relative noise level (solid line) and the sky transmission (shaded area). The red solid line shows the best-fit model (see Section 3.1). Positions of major emission and absorption lines are indicated by dot-dashed (blue) and dashed (red) lines, respectively, even when not detected.
(A color version of this figure is available in the online journal.)

(Cimatti et al. 2008). Cappellari et al. (2009) found two galaxies with similar stellar density and $\sigma_\star$ as local ETGs of the same mass. The remaining galaxies have higher stellar densities and higher $\sigma_\star$ from their stacked spectrum, but still overlapping with the densest local ETGs.

In this respect, near-infrared (NIR) spectroscopy offers a great advantage for the galaxies at $z \gtrsim 1.4$ as the optical break at rest frame 4000 Å and Ca II H+K (the strongest spectral features of passively evolving galaxies) are redshifted into NIR, and the rest-frame optical continuum is much brighter than that in the rest-frame UV. Thus, van Dokkum et al. (2009) measured a velocity dispersion of $\sigma_\star \simeq 500$ km s^{-1} for a compact ETG at $z = 2.186$ using a deep NIR spectrum from Kriek et al. (2009). This value of $\sigma_\star$ is much higher than that of the most massive local ETGs, and would be consistent with the small half light radius measured for that galaxy.

In this Letter, we present a rest-frame optical spectrum of a massive, passively evolving high-redshift galaxy candidate taken with the Multi-Object InfraRed Camera and Spectrograph (MOIRCS; Ichikawa et al. 2006; Suzuki et al. 2008) at the Subaru Telescope. A cosmology with $\Omega_{\rm M} = 0.3$, $\Omega_\Lambda = 0.7$, and $H_0 = 70$ km s^{-1} Mpc^{-1} is assumed.

2. OBSERVATIONS AND DATA REDUCTIONS

We have obtained 4.7 hr of MOIRCS spectroscopy of 34 BzK-selected galaxies (Daddi et al. 2004) from the catalog of McCracken et al. (2010) in the COSMOS field. We used the zJ500 grism with $0\farcs7$ slits, providing a resolution of $R \simeq 500$ in the J band over the range 9500–16000 Å. The primary aim of our observation was to measure redshifts for passive BzK galaxies (pBzKs) and to locate them accurately in the COSMOS large-scale structure. We preferentially selected the most massive pBzKs (which are also the brightest in the NIR) to maximize the chance of determining also physical information in addition to redshifts for at least a fraction of them, in particular the galaxy (254025) discussed in this Letter and already studied by Mancini et al. (2010). The observations were made under partly cloudy conditions and with $\sim 1\farcs2$ seeing. A sequence of 600 s integrations were made in a two-position dithering pattern separated by $2''$. The A0V-type star HIP 55627 was observed to obtain (relative) flux calibration and to correct for telluric absorption. The data were reduced using the *MCSMDP* pipeline (Yoshikawa et al. 2010), including flat-fielding by dome flat, sky subtraction between each exposure pair, bad pixel and cosmic-ray rejection, distortion correction, wavelength calibration (based on the OH telluric lines), residual sky subtraction, and finally co-addition with appropriate offsets and weights. The two-dimensional spectra are flux-calibrated using the standard star spectrum, and one-dimensional spectra were extracted with the IRAF *apall* task using a $1\farcs9$ aperture. The absolute flux calibration was then obtained by normalizing to the J-band total magnitude.

While the results for the complete sample observed with MOIRCS will be presented elsewhere, we will concentrate here on the pBzK galaxy 254025. This galaxy is one of the 12 ultra-massive high-redshift ETG candidates in Mancini et al. (2010), with a photometric redshift of $z_{\rm phot} = 1.71$ and very bright NIR magnitudes of $J_{\rm AB} = 20.32$ and $K_{\rm AB} = 19.41$. Mancini et al. (2010) also report that the galaxy is non-detected at *Spitzer*/MIPS 24 μm to 80 μJy implying a star formation rate (SFR) $\lesssim 50\,M_\odot$ yr^{-1}. Using *HST*/ACS F814W image (rest-frame UV) they measure a Sérsic index $n = 4.1$ and an effective radius of $r_{\rm e} = 5.7$ kpc, consistent with the stellar mass-size relation of local elliptical galaxies (Mancini et al. 2010).

Figure 1 shows the resulting one- and two-dimensional MOIRCS spectra of 254025. The 4000 Å break is clearly seen, together with strong Balmer and metallic absorption lines, namely, Ca II H +Hϵ, Ca II K, Hδ, Hγ, Hζ, G band, and CN+Hθ. Hβ falls in a region with low atmospheric transmission and with strong OH lines. No emission lines are observed. While [O III]$\lambda\lambda$4959, 5007 falls in a region with low atmospheric transmission, [O II]λ3727 is uncontaminated and its non-detection sets a 3σ upper limit of an SFR of $\simeq 2.5\,M_\odot$ yr^{-1} (not corrected for extinction), using the Kennicutt (1998) conversion.

Table 1
Best-fit Stellar Population Parameters for the SED and the Spectrum

SFH	SED						Spectrum					
	$T_{\rm onset}$ (Gyr)	τ or t_q (Gyr)	$t_{\rm sf}$ (Gyr)	$M_\star$ ($10^{11} M_\odot$)	M/L_U ($(M/L_U)_\odot$)	χ^2	$T_{\rm onset}$ (Gyr)	τ or t_q (Gyr)	$t_{\rm sf}$ (Gyr)	$M_\star$ ($10^{11} M_\odot$)	M/L_U ($(M/L_U)_\odot$)	χ^2
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)
$Z = 0.5Z_\odot$												
All	···	···	$2.54^{+0.24}_{-0.16}$	$4.84^{+0.58}_{-0.30}$	$1.11^{+0.11}_{-0.07}$	2.35	···	···	$1.16^{+0.86}_{-0.01}$	$2.33^{+1.10}_{-0.02}$	$0.45^{+0.21}_{-0.01}$	1.34
SSP	0.7	···	$0.70^{+0.01}_{-0.01}$	$1.58^{+0.01}_{-0.01}$	$0.35^{+0.01}_{-0.01}$	9.25	1.2	···	$1.20^{+0.01}_{-0.03}$	$2.40^{+0.01}_{-0.01}$	$0.47^{+0.01}_{-0.01}$	1.35
Const. + quenching	3.4	3.00	$1.90^{+0.03}_{-0.61}$	$2.83^{+0.16}_{-0.67}$	$0.64^{+0.02}_{-0.15}$	3.99	1.2	0.09	$1.16^{+0.86}_{-0.04}$	$2.33^{+1.24}_{-0.03}$	$0.45^{+0.23}_{-0.01}$	1.34
Delayed exponential	3.2	0.33	$2.54^{+0.24}_{-0.56}$	$4.84^{+0.58}_{-1.20}$	$1.11^{+0.11}_{-0.24}$	2.35	1.2	0.02	$1.16^{+0.13}_{-0.04}$	$2.32^{+0.25}_{-0.04}$	$0.45^{+0.04}_{-0.01}$	1.34
Exp. + quenching	3.4	2.87	$1.20^{+0.03}_{-0.36}$	$1.98^{+0.10}_{-0.33}$	$0.45^{+0.02}_{-0.08}$	6.84	1.2	0.10	$1.15^{+0.39}_{-0.02}$	$2.32^{+0.61}_{-0.03}$	$0.45^{+0.11}_{-0.01}$	1.34
$Z = Z_\odot$												
All	···	···	$1.14^{+0.73}_{-0.01}$	$2.76^{+0.82}_{-0.01}$	$0.63^{+0.21}_{-0.01}$	2.20	···	···	$1.95^{+0.01}_{-0.09}$	$3.91^{+0.07}_{-0.16}$	$0.77^{+0.02}_{-0.04}$	1.31
SSP	0.6	···	$0.60^{+0.01}_{-0.01}$	$1.62^{+0.01}_{-0.01}$	$0.34^{+0.01}_{-0.01}$	16.22	1.0	···	$1.00^{+0.01}_{-0.01}$	$2.75^{+0.02}_{-0.03}$	$0.54^{+0.01}_{-0.01}$	1.33
Const. + quenching	3.4	3.05	$1.88^{+0.02}_{-0.55}$	$3.36^{+0.11}_{-0.75}$	$0.75^{+0.02}_{-0.17}$	2.28	3.5	3.11	$1.95^{+0.02}_{-0.92}$	$3.91^{+0.16}_{-1.19}$	$0.77^{+0.04}_{-0.23}$	1.31
Delayed exponential	1.4	0.13	$1.14^{+0.70}_{-0.08}$	$2.76^{+1.52}_{-0.24}$	$0.63^{+0.40}_{-0.05}$	2.20	1.2	0.07	$1.06^{+0.05}_{-0.04}$	$2.78^{+0.12}_{-0.14}$	$0.54^{+0.02}_{-0.03}$	1.32
Exp. + quenching	1.7	1.15	$0.98^{+0.17}_{-0.06}$	$2.19^{+0.16}_{-0.16}$	$0.48^{+0.03}_{-0.03}$	6.62	1.2	0.31	$1.03^{+0.38}_{-0.04}$	$2.74^{+0.62}_{-0.11}$	$0.54^{+0.12}_{-0.02}$	1.32
$Z = 2Z_\odot$												
All	···	···	$0.97^{+0.86}_{-0.27}$	$2.22^{+1.29}_{-0.26}$	$0.52^{+0.34}_{-0.07}$	4.32	···	···	$1.88^{+0.01}_{-0.24}$	$3.99^{+0.10}_{-0.32}$	$0.80^{+0.03}_{-0.07}$	1.29
SSP	0.4	···	$0.40^{+0.01}_{-0.01}$	$1.43^{+0.01}_{-0.01}$	$0.31^{+0.01}_{-0.01}$	34.92	0.8	···	$0.80^{+0.02}_{-0.01}$	$2.65^{+0.09}_{-0.04}$	$0.53^{+0.02}_{-0.01}$	1.31
Const. + quenching	1.8	1.62	$0.99^{+0.85}_{-0.45}$	$2.35^{+1.25}_{-0.70}$	$0.55^{+0.32}_{-0.18}$	4.46	3.5	3.25	$1.88^{+0.01}_{-0.83}$	$3.99^{+0.16}_{-1.07}$	$0.80^{+0.04}_{-0.22}$	1.29
Delayed exponential	1.0	0.10	$0.80^{+0.01}_{-0.01}$	$2.39^{+0.01}_{-0.01}$	$0.61^{+0.01}_{-0.01}$	9.39	1.1	0.10	$0.90^{+0.03}_{-0.09}$	$2.78^{+0.12}_{-0.17}$	$0.55^{+0.03}_{-0.03}$	1.31
Exp. + quenching	2.9	2.58	$0.97^{+0.05}_{-0.32}$	$2.22^{+0.17}_{-0.46}$	$0.52^{+0.04}_{-0.12}$	4.32	3.3	3.01	$1.80^{+0.02}_{-0.81}$	$3.95^{+0.15}_{-1.11}$	$0.79^{+0.04}_{-0.22}$	1.29

Notes. Column 1: SFH (see Section 3.1); Columns 2 and 8: elapsed time since the onset of SF; Columns 3 and 9: SF timescale in the case of delayed exponential SFH and quenching time for SFHs with constant SFR+quenching and exponentially increasing SFR+quenching; Columns 4 and 10: SF-weighted age defined by $\int_0^T (T-t)\phi(t)\,dt / \int_0^T \phi(t)\,dt$, where T is $T_{\rm onset}$ and $\phi(t)$ is SFR; Columns 5 and 11: stellar mass; Columns 6 and 12: rest-frame U-band mass-to-light ratio; Columns 7 and 13: reduced χ^2 for the best-fit template.

From the spectrum, the absorption line redshift is measured as $z = 1.8230 \pm 0.0006$.

3. RESULTS

3.1. Stellar Populations

Having determined the spectroscopic redshift, we proceeded to fit stellar population templates, separately to the broadband spectral energy distribution (SED) and to the MOIRCS spectrum. We allowed for a wide range of possible star formation histories (SFHs), including: (1) instantaneous bursts, i.e., simple stellar populations (SSPs); (2) constant SFRs for a duration within 0.01–3.5 Gyr,[19] terminated by SF quenching and followed by passive evolution; (3) delayed, exponentially declining SFH described as $\mathrm{SFR}(t,\tau) \propto (t/\tau^2)\exp(-t/\tau)$ with τ within 0.01–2 Gyr; and (4) exponentially increasing SFH, $\mathrm{SFR}(t,\tau) \propto \exp(t/\tau)$ for a duration within $t_q = 0.1$–3 Gyr, followed by SF quenching and passive evolution. We choose $\tau = 0.72$ Gyr, corresponding to a stellar mass doubling time of $\simeq$0.5 Gyr, as suggested for $z \sim 2$ galaxies by the existence of tight stellar-mass–SFR relation, with $\mathrm{SFR} \propto \sim M_\star$ (Daddi et al. 2007; Renzini 2009). For all the SFHs, template ages were allowed to range in $t = 0.4$–3.5 Gyr. We use a Chabrier (2003) initial mass function. The fits were made with metallicities of $0.5 \times Z_\odot$, $Z_\odot$ and $2 \times Z_\odot$. To reduce the number of free parameters, we have assumed no dust extinction, appropriate for a passively evolving galaxy, considering the strict upper limit on the SFR that is set by the spectrum, an assumption that is validated by the good fit that is achieved in the blue continuum (see Figure 1).

The SED fitting was carried out for the broadband Biz data from Subaru/Suprime-Cam (Capak et al. 2007; Taniguchi et al. 2007), JHK data from CFHT/WIRCAM (McCracken et al. 2010), and the *Spitzer*/IRAC 3.6 μm, 4.5 μm, and 5.8 μm data (Sanders et al. 2007). Artificial errors of 0.05 mag for the $BizJHK$ bands and 0.1 mag for the IRAC bands are added to the observed errors in quadrature to account for systematics in zero-point determinations, in the photometric measurements, and in the stellar population models. The templates for the SED fitting are generated from population synthesis models by Maraston (2005, hereafter M05).

The spectral resolution of the M05 models is significantly lower than that of our MOIRCS spectrum. Hence, for fitting the spectrum, we used templates from the Bruzual & Charlot (2003, hereafter BC03) spectral synthesis library. Although these models might not account properly for TP-AGB stars (e.g., Maraston et al. 2006), this effect is not significant at the 4000 Å rest frame, the wavelength range probed by our spectrum. The template spectra are Gaussian-broadened to an overall velocity dispersion of 350 km s^{-1} (see Section 3.2) to match the observed spectrum, and having fixed it the stellar population parameters are derived with the χ^2 over the observed wavelength range 9500–16000 Å.

The stellar population parameters of the best-fit models from each adopted SFH are listed in Table 1, and the best-fit templates for the spectrum and SED are shown in Figures 1 and 2, respectively. The best-fit models have SF-weighted ages of

[19] The age of the universe at $z = 1.82$ is about 3.5 Gyr, given the adopted cosmology.

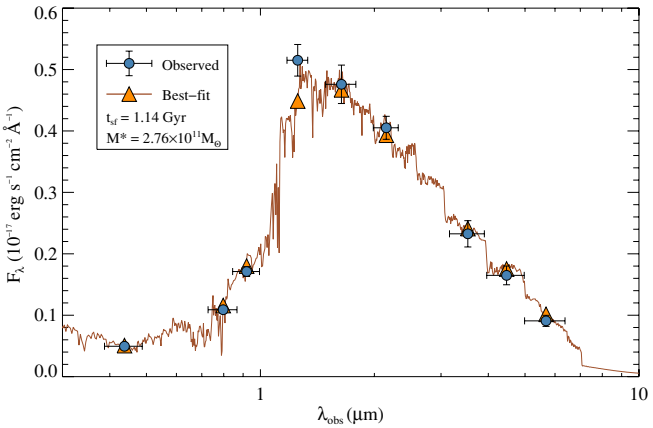

Figure 2. Observed SED for 254025 (blue circles with error bars) compared to the best-fit model from Maraston (2005; orange line and symbols). The best fit was found for an SFH with a delayed exponential SFR that continued for about 1.4 Gyr, with an SF timescale of 0.13 Gyr, and for a solar metallicity. The parameters of the best-fit templates can be found in Table 1.

(A color version of this figure is available in the online journal.)

$t_{\rm sf} = 1.14^{+0.73}_{-0.01}$ Gyr and stellar masses of $M_\star = (2.76^{+0.82}_{-0.01}) \times 10^{11}\,M_\odot$ for the broadband SED, and $t_{\rm sf} = 1.88^{+0.01}_{-0.24}$ Gyr and stellar masses of $M_\star = (3.99^{+0.10}_{-0.32}) \times 10^{11}\,M_\odot$ for the spectrum. The best-fit results are from $Z = Z_\odot$ for the SED and $Z = 2 \times Z_\odot$ for the spectrum. However, for the spectrum very similar values are derived using solar metallicity models that result in a slightly higher χ^2 (Table 1). We note that solar, or slightly supersolar metallicities are appropriate for local elliptical galaxies with stellar masses similar to galaxy 254025 (Thomas et al. 2005). The SF-weighted age of $\simeq$1–2 Gyr corresponds to an average formation redshift of $z_{\rm form} \simeq 2.5$–4, although the SF could have started much earlier. In the case of the spectral fitting, all SFHs adopted here fit equally well with $\chi^2 \simeq 1.3$ and $t_{\rm sf} \simeq$ 1–2 Gyr, consistent with the detection of strong Balmer absorption lines which are most prominent in A-type stars. Moreover, in the case of the SED fits (apart from the SSP spectra) the various SFHs do not give very different χ^2 values and therefore it is not possible to discriminate between them. The same can be said for the derived metallicities. If we allow for dust reddening, the best fitting $t_{\rm sf}$, $M_\star$, and M/L_U would change by only $\lesssim$10%, with some increase in the formal uncertainties within each class of SFH.

3.2. Velocity Dispersion and Dynamical Modeling

Our high signal-to-noise ratio (S/N) spectrum ($\simeq$8.7 per 60 km s^{-1} spectral interval in the continuum) allows us to measure a velocity dispersion from the absorption line width $\sigma_{\rm obs}$, which is a combination of the galaxy stellar velocity dispersion $\sigma_\star$ and the instrumental resolution $\sigma_{\rm instr}$. This S/N is comparable to spectra of GMASS galaxies with successful individual $\sigma_\star$ determinations (Cappellari et al. 2009). Therefore, we followed the same approach of Cappellari et al., based on the Penalized Pixel-Fitting method (pPXF; Cappellari & Emsellem 2004). The MILES stellar library containing 985 stars (Sánchez-Blázquez et al. 2006) is adopted here since it provides the best uniform and complete set of stars.

Figure 3 shows the best-fit templates from pPXF, corresponding to $\sigma_{\rm obs} = 350 \pm 30$ km s^{-1} (random) $\pm$ 30 km s^{-1} (systematic) for the rest-frame wavelength range of 3500–4450 Å. The random error (1σ confidence) is determined as half of the interval in $\sigma_{\rm obs}$ spanned by 68 out of 100 Monte Carlo realizations of the input spectrum. A rough estimate of the systematic error is obtained as half of the interval in $\sigma_{\rm obs}$ spanned by all repeated extractions of the kinematics using different, but equally acceptable, combinations for the values of the degree (from 0–4) of the additive and multiplicative polynomials in pPXF. Restricting the fit to the region with the Balmer and Ca II H+K lines (3700–4100 Å in the rest frame) gives $\sigma_{\rm obs} = 300 \pm 50$ km s^{-1} (total error), consistent with the value derived from the full spectral range.

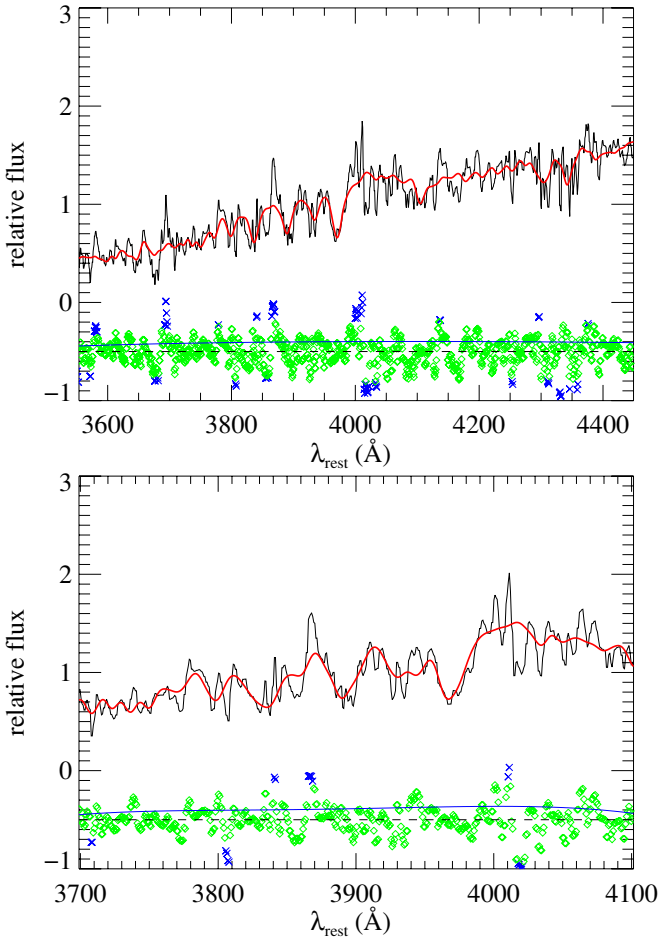

Figure 3. Results of the pPXF fit for the stellar velocity dispersion of the galaxy. The panels show the resulting fit for the full spectral range (top) and for a wavelength range around Ca II H+K (bottom). The black solid line shows the observed spectrum; the red solid line shows the best-fit template; the green diamonds are the residuals (arbitrarily shifted). The blue crosses indicate bad pixels rejected from the fitting. The solid blue line indicates the estimated 1σ noise level.

(A color version of this figure is available in the online journal.)

In order to derive $\sigma_\star$, we need to determine accurately the instrumental resolution. To do this, we have used our combined MOIRCS spectrum without sky subtraction and simultaneously fitted Gaussian profiles to a series of telluric OH lines at $\lambda \simeq 11500$ Å, i.e., near the strongest absorption features of the galaxy's spectrum. The central wavelength for each OH line was taken from Rousselot et al. (2000) and we left $\sigma_{\rm instr}$, OH-line intensities, and constant baseline as free parameters. The fitting procedure reproduces the observed sky spectra very well. Over $\lambda = 3500$–4450 Å, the instrumental resolution changes from 270 km s^{-1} to 330 km s^{-1}. We adopt $\sigma_{\rm instr} = 300 \pm 7$ km s^{-1} (random) $\pm$30 km s^{-1} (systematic).

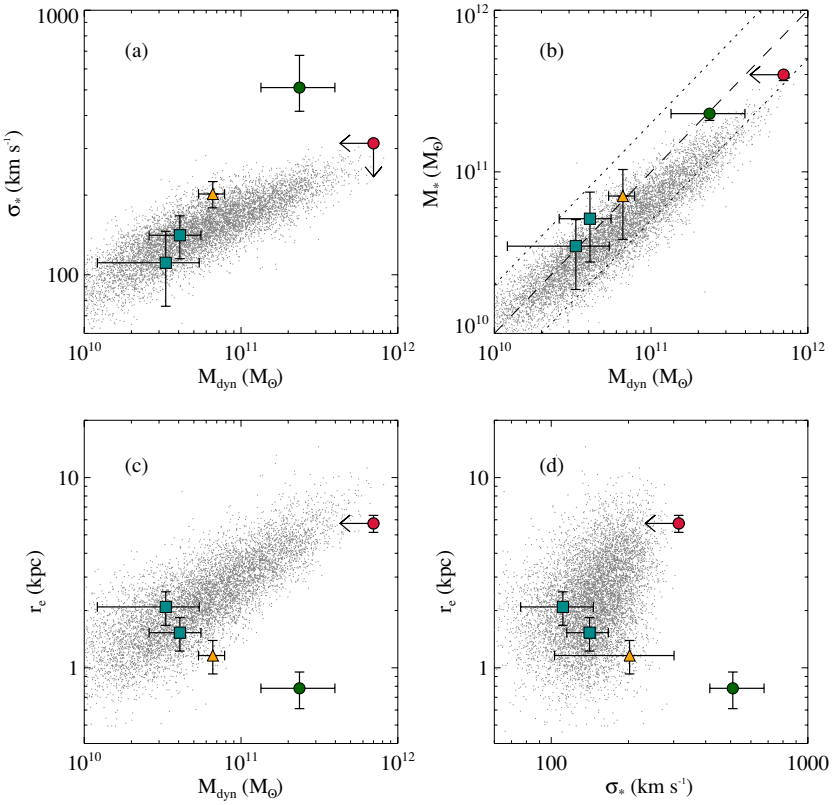

Figure 4. Comparison of the properties of high-redshift ETGs for which the velocity dispersion has been measured so far (symbols with error bars) with those of elliptical galaxies selected from SDSS at $z \simeq 0.06$ (gray dots). (a) the stellar velocity dispersions vs. the virial masses. (b) comparison between virial and stellar masses; the diagonal dashed line corresponds to equality of the two masses, and the dotted lines show a range by a factor of 2. (c) the effective radii vs. the virial masses. (d) the effective radii vs. stellar velocity dispersions. In all panels, the red filled circle represents the galaxy studied here; the green filled circle shows the object studied by van Dokkum et al. (2009); the blue squares represent the two GMASS galaxies with individual $\sigma_\star$ measurements; the yellow triangle represents the properties from the stacked GMASS spectrum (taken from Cappellari et al. 2009).

(A color version of this figure is available in the online journal.)

The derived galaxy stellar velocity dispersion is $\sigma_\star = \sqrt{\sigma^2_{\rm obs} - \sigma^2_{\rm instr}}$, which gives $\sigma_\star = 180 \pm 59$ km s^{-1} (random) ± 87 km s^{-1} (systematic). Due to relatively large uncertainties in $\sigma_{\rm obs}$, and as $\sigma_{\rm obs}$ is close to $\sigma_{\rm instr}$, we cannot place a lower limit to $\sigma_\star$. However, we can derive a 1σ upper limit of $\sigma_\star < 326$ km s^{-1} (or $\sigma_\star < 385$ km s^{-1} at the 2σ level), which is consistent with both determinations. If Balmer lines suffer from fill-in from emission lines $\sigma_\star$ could be somewhat lower.

From the stellar velocity dispersion the virial mass can be calculated as $M_{\rm vir} = C r_{\rm e} \sigma^2_\star / G$. We have set $C = 5$ as empirically calibrated on local galaxies with state-of-the-art dynamical modeling (Cappellari et al. 2006), with the velocity dispersion being measured within a large aperture ($\sim 1 r_{\rm e}$) as in our case. The effective radius $r_{\rm e}$ was measured by Mancini et al. (2010) from the *HST*/ACS F814W image ($\simeq$2900 Å in the rest frame) as $0\farcs68 \pm 0\farcs07$ or 5.7 ± 0.6 kpc at $z = 1.82$. Thus, the upper limit of the virial mass is derived as $M_{\rm vir} < 7.0 \times 10^{11}\, M_\odot$.

We have also constructed a dynamical model based on axisymmetric Jeans dynamical models as those used to model the GMASS galaxies by Cappellari et al. (2009), adopting a multi-Gaussian expansion (MGE; Emsellem et al. 1994). This method has the advantage that the derived M/L is virtually insensitive to possible underestimation of the size, which can be a possibility at high redshifts. Considering the bolometric surface brightness dimming of $(1+z)^4$, a factor $(1+z)$ coming from the source redshifting, and the K-correction between rest frame 2900 Å and rest-frame U band, we derived a rest-frame U-band luminosity of $L_U = 6.7 \times 10^{11}\, L_\odot$. The second moment of the velocity $V^2_{\rm rms} = V^2 + \sigma^2_\star$ was also estimated (assuming $\beta_z = 0$ and axisymmetry) by using the Jeans anisotropic MGE (JAM) method (Cappellari 2008). The upper limit for the dynamical M/L_U can be calculated by $(M/L_U)_{\rm Jeans} = (\sigma_\star / V_{\rm rms})^2 < 1.0$, which can be converted into the upper limit of the dynamical mass from the JAM model as $M_{\rm Jeans} = L_U \times (M/L_U)_{\rm Jeans} < 6.8 \times 10^{11}\, M_\odot$. Therefore, the virial mass and Jeans mass agree well though both of them are upper limits. A JAM model constructed from a noiseless model with the best-fitting Sérsic parameters of Mancini et al. (2010), as opposed to the actual Advanced Camera for Surveys (ACS) image, gives the same $(M/L_U)_{\rm Jeans}$ within 1%. This is due to the robustness of the central M/L recovered using dynamical models (in contrast to virial estimates) to photometric uncertainties at large radii (e.g., Section 3.2 in Cappellari et al. 2009).

4. DISCUSSION AND CONCLUSIONS

Figure 4 compares the properties of the galaxy 254025 and other $z \simeq 2$ galaxies for which the same quantities have been measured (Cappellari et al. 2009; van Dokkum et al. 2009). The figure includes ETGs at $z \simeq 0.06$, selected from the Sloan Digital Sky Survey (SDSS) on the basis of their red $u - g$ color and high Sérsic index $n \simeq 4$ (Blanton et al. 2005). The dynamical and stellar masses agree very well for the high-z objects, within a factor of $\lesssim$2. Note that our massive galaxy has physical properties in good agreement with those of local ETGs

of similar stellar mass. Our galaxy provides a second example of a very massive passively evolving system for which a stellar velocity dispersion has been measured (the GMASS objects of Cappellari et al., which in Figure 4 also lie on the $z = 0$ scaling relations, but have stellar masses below $10^{11}\ M_{\odot}$). The "normal" size and velocity dispersion of our massive ETG is strikingly in contrast to the extreme properties (i.e., a very high $\sigma_{\star} = 510^{+163}_{-95}$ km s^{-1} and a small $r_{\rm e} = 0.78 \pm 0.17$ kpc) of the galaxy studied by van Dokkum et al. (2009) with similar stellar mass ($2 \times 10^{11}\ M_{\odot}$). This suggests a substantial diversity in the physical properties of the *massive* ETG population at $z \sim 2$ including "immature," albeit virialized, systems—which will have to evolve into normal $z = 0$ massive galaxies through some physical processes which decrease their velocity dispersion and increase their sizes—as well as "mature" ETGs, already on the scaling relationships of $z = 0$ ETGs. It is clear that many more observations of similar galaxies are required to establish which kind of ETG is commonest at high redshift: either the compact/high-$\sigma_{\star}$ objects like those found by van Dokkum et al. (2009), or the apparently normal, low-$\sigma_{\star}$ objects presented in this Letter. Also, nothing prevents our particular object to evolve further from its present state which mimics that of local ellipticals of the same mass. For example, it may grow further and become a brightest cluster galaxy (BCG) or a cD galaxy. For this reason, it would be important to estimate the volume number density of similar objects at high redshifts and to compare it to that of BCGs and cD galaxies.

To conclude, both very compact ETGs and ETGs following the local stellar-mass–size and stellar-mass–$\sigma_{\star}$ relations appear to co-exist at $z > 1.4$ (see also Mancini et al. 2010). However the number of high-z ETGs with individual measurement of the velocity dispersion is still extremely small. Increasing their sample is of great importance to understand the evolution of these galaxies, and in particular how and when they acquire their final structural and dynamical configuration. This Letter demonstrates that with reasonable telescope time several absorption features can be detected in the rest-frame optical spectrum of the high-z ETGs, from which (at least for the most massive ETGs) the velocity dispersion and several stellar population properties can be derived.

We are grateful to Tomohiro Yoshikawa for providing *MCSMDP* before publication. We thank the Subaru telescope staff for help with our observations. We acknowledge funding ERC-StG-UPGAL-240039, ANR-07-BLAN-0228, ANR-08-JCJC-0008 and a Grant-in-Aid for Science Research (No. 19540245) by the Japanese Ministry of Education, Culture, Sports, Science, and Technology. A.R. is grateful to the Institute of Astronomy of ETH Zürich for its kind hospitality. M.C. acknowledges support from an STFC Advanced Fellowship (PP/D005574/1).

REFERENCES

Blanton, M. R., et al. 2005, AJ, 129, 2562
Bruzual, G., & Charlot, S. 2003, MNRAS, 344, 1000
Capak, P., et al. 2007, ApJS, 172, 99
Cappellari, M. 2008, MNRAS, 390, 71
Cappellari, M., & Emsellem, E. 2004, PASP, 116, 138
Cappellari, M., et al. 2006, MNRAS, 366, 1126
Cappellari, M., et al. 2009, ApJ, 704, L34
Cenarro, A. J., & Trujillo, I. 2009, ApJ, 696, L43
Chabrier, G. 2003, PASP, 115, 763
Cimatti, A., et al. 2004, Nature, 430, 184
Cimatti, A., et al. 2008, A&A, 482, 21
Daddi, E., Cimatti, A., Renzini, A., Fontana, A., Mignoli, M., Pozzetti, L., Tozzi, P., & Zamorani, G. 2004, ApJ, 617, 746
Daddi, E., et al. 2005, ApJ, 626, 680
Daddi, E., et al. 2007, ApJ, 670, 156
Emsellem, E., Monnet, G., & Bacon, R. 1994, A&A, 285, 723
Fan, L., Lapi, A., De Zotti, G., & Danese, L. 2008, ApJ, 689, L101
Feldmann, R., Carollo, C. M., Mayer, L., Renzini, A., Lake, G., Quinn, T., Stinson, G. S., & Yepes, G. 2010, ApJ, 709, 218
Hopkins, P. F., Bundy, K., Hernquist, L., Wuyts, S., & Cox, T. J. 2010, MNRAS, 401, 1099
Ichikawa, T., et al. 2006, Proc. SPIE, 6269, 38
Kennicutt, R. C., Jr. 1998, ARA&A, 36, 189
Khochfar, S., & Silk, J. 2006, ApJ, 648, L21
Kriek, M., van Dokkum, P. G., Labbé, I., Franx, M., Illingworth, G. D., Marchesini, D., & Quadri, R. F. 2009, ApJ, 700, 221
La Barbera, F., & de Carvalho, R. R. 2009, ApJ, 699, L76
Longhetti, M., et al. 2007, MNRAS, 374, 614
Mancini, C., et al. 2010, MNRAS, 401, 933
Maraston, C. 2005, MNRAS, 362, 799
Maraston, C., Daddi, E., Renzini, A., Cimatti, A., Dickinson, M., Papovich, C., Pasquali, A., & Pirzkal, N. 2006, ApJ, 652, 85
McCarthy, P. J., et al. 2004, ApJ, 614, L9
McCracken, H. J., et al. 2010, ApJ, 708, 202
Naab, T., Johansson, P. H., & Ostriker, J. P. 2009, ApJ, 699, L178
Pannella, M., et al. 2009, ApJ, 698, L116
Renzini, A. 2009, MNRAS, 398, L58
Rousselot, P., Lidman, C., Cuby, J., Moreels, G., & Monnet, G. 2000, A&A, 354, 1134
Sánchez-Blázquez, P., et al. 2006, MNRAS, 371, 703
Sanders, D. B., et al. 2007, ApJS, 172, 86
Saracco, P., Longhetti, M., & Andreon, S. 2009, MNRAS, 392, 718
Suzuki, R., et al. 2008, PASJ, 60, 1347
Taniguchi, Y., et al. 2007, ApJS, 172, 9
Thomas, D., Maraston, C., Bender, R., & Mendes de Oliveira, C. 2005, ApJ, 621, 673
Trujillo, I., et al. 2006, ApJ, 650, 18
van Dokkum, P. G., Kriek, M., & Franx, M. 2009, Nature, 460, 717
van Dokkum, P. G., et al. 2008, ApJ, 677, L5
Yoshikawa, T., et al. 2010, ApJ, submitted